Ionizing Radiation and Life

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Abstract

Ionizing radiation is a ubiquitous feature of the Cosmos, from exogenous cosmic rays (CR) to the intrinsic mineral radioactivity of a habitable world, and its influences on the emergence and persistence of life are wide-ranging and profound. Much attention has already been focused on the deleterious effects of ionizing radiation on organisms and the complex molecules of life, but ionizing radiation also performs many crucial functions in the generation of habitable planetary environments and the origins of life. This review surveys the role of CR and mineral radioactivity in star formation, generation of biogenic elements, and the synthesis of organic molecules and driving of prebiotic chemistry. Another major theme is the multiple layers of shielding of planetary surfaces from the flux of cosmic radiation and the various effects on a biosphere of violent but rare astrophysical events such as supernovae and gamma-ray bursts. The influences of CR can also be duplicitous, such as limiting the survival of surface life on Mars while potentially supporting a subsurface biosphere in the ocean of Europa. This review highlights the common thread that ionizing radiation forms between the disparate component disciplines of astrobiology. Key Words: Cosmic rays—Solar energetic particles—Radioactivity—Ultraviolet—Supernova—Gamma-ray burst—Magnetosphere—Panspermia—Biosignature. Astrobiology 11, 551–582.

1. Introduction

The surface of Earth is a relatively benign environment with regard to ionizing particle radiation. The two major natural sources of ionizing radiation on a planetary body are the decay of radioactive isotopes present on the surface and the flux of cosmic rays from above. Cosmic rays (CR) are charged subatomic particles that originate in space and are accelerated to high energies. Cosmic rays are a ubiquitous source of ionizing radiation throughout the Cosmos that impacts habitable worlds and the origin and maintenance of life in a plethora of different ways. This energetic particle radiation has been studied as a hazard of human space exploration or, more generally, as a radiation threat to organisms in the space environment or on a planetary surface (e.g., see recent reviews by Brack et al., 2010, and Horneck et al., 2010); but the influences of CR and planetary radioactivity upon life extend far beyond their direct radiobiological hazard. Ionizing radiation performs many crucial functions in astrobiology, including its role in star formation, driving astrochemistry and the synthesis of extraterrestrial organic molecules, and perhaps supporting a subsurface europaan biosphere. The high flux of energetic electromagnetic radiation from astrophysical events such as gamma-ray bursts (GRB) is also considered in this review, and though radiation from these types of events is not termed CR, it can cause extreme ionization perturbations and seriously threaten planetary habitability.

Other authors have provided recent overviews of different aspects of ionizing radiation and astrobiology, and interested readers may like to consult some of these for further discussion. A study by Ferrari and Szuszkiewicz (2009) offers an excellent introduction for nonspecialists to the sources and interactions of CR and details of biological dosimetry. For reviews that focus on the nature of CR and their interaction with shielding matter see works by Amsler et al. (2008a, 2008b). Baumstark-Khan and Facius (2001) provided a very useful account of different sources of ionizing radiation, including radionuclide decay, and their biological effects, while Nelson (2003) focused on the cellular effects of CR strikes. Vainio et al. (2009) summarized the dynamics of the particle radiation environment at Earth, whereas Bazilevskaya et al. (2008) provided a more focused account of the ionization of Earth’s atmosphere; and Bailer-Jones (2009) critically assessed claims that CR flux affects the terrestrial climate. Lammer et al. (2009) offered a broad overview of the influences on planetary habitability, including that of CR. Finally, the published conference proceedings by Akaboshi et al. (2000) present a selection of papers that address the roles of radiation in the origin and evolution of life.

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This paper reviews the current state of knowledge on the different roles of ionizing radiation in the emergence and maintenance of life, and so argues that CR act as a common thread between many of the disparate areas within astrobiology. Ionizing radiation is one of the major factors that reaches across the disciplines of astrobiology—linking life in a protected planetary environment to its cosmic context.

The review will not deal extensively with UV light, such as that produced by the broad spectrum stellar emission. Here, I am primarily concerned with the ionizing radiation environment produced by CR, as well as extremely energetic electromagnetic radiation from astrophysical sources such as GRB. In addition, I discuss the role of the ionizing radiation field intrinsic to a habitable world generated by radionuclide decay in its crust or oceans where it is relevant to the origin and survival of life.

This review is not intended as a comprehensive or definitive resource and does not attempt to collate all related papers over the last few decades in fields as diverse as astrophysics, prebiotic chemistry, and planetary habitability. It is written as a broad overview for nonspecialists—an introductory resource to allow the reader to appreciate the wide-ranging influences that ionizing radiation exerts in astrobiology and to provide some key papers for follow-up reading.

I will begin by explaining the astrophysical sources of CR and their deleterious effect on living organisms. I will then progress through the greatly diverse range of topics in astrobiology unified by CR, explaining the role of CR in star formation, in the generation of elements needed by life, and in driving astrochemistry and the production of organic molecules. Cosmic rays also play a major role in the habitability of a world, and I will review how a planetary surface is shielded by the star's astrosphere, the planet's magnetosphere (if present) and atmosphere, and the influences of CR in forcing the climate or atmospheric chemistry. Cosmic ray ionization in the atmosphere may play an important role in the astrobiological potential of Venus and Titan, as well as on unshielded planetary bodies such as Mars, Europa, and meteorites, in the interest of pansperma. The next section deals with extreme, but irregular, ionizing radiation hazards, such as supernovae, GRB, and superflares, and the predicted effect on life, such as destruction of the ozone layer, global cooling, nitric acid rain-out, and intense UV pulse. Finally, I will review the possible roles of intrinsic radioactivity in planetary habitability and prebiotic chemistry and the occurrence of natural fission reactors. The structure of this review is as follows:

2. Cosmic Ray Sources
3. Biological Effects of Ionizing Radiation
4. The Role of Cosmic Rays in Star Formation
5. Role of Cosmic Rays in Biogenic Elements
6. Role of Cosmic Rays in Astrochemistry and Prebiotic Organic Molecules
7. Influence of Cosmic Rays on Planetary Habitability
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   7.2. Planetary magnetosphere
   7.3. Interaction of cosmic rays with atmosphere
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9. Intrinsic Planetary Radiation Environment
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   9.2. Natural fission reactors
10. Conclusions

2. Cosmic Ray Sources

Beyond the shielding influence of Earth's magnetic field and atmosphere (see below), the space radiation environment is dominated by solar energetic particles (SEP; but also referred to as solar cosmic rays, SCR) and galactic cosmic rays (GCR), as illustrated in Fig. 1. SEP and GCR are summarized briefly here to provide context to the astrobiological effects of ionizing radiation discussed later; more complete reviews of CR can be found in studies by Ferrari and Szuszkiewicz (2009), Vainio et al. (2009), and Amsler et al. (2008a, 2008b).

Solar energetic particles, primarily protons, are accelerated by flares and coronal mass ejections to energies typically of tens to hundreds of megaelectronvolts. Particularly large particle events release an appreciable fluence up to around 1 GeV and cause significant ground level enhancements in CR intensity in detectors on Earth's surface (SEP were reviewed by Ryan et al., 2000, and acceleration mechanisms were reviewed recently by Giacalone and Kóta, 2006). The flux of these SEP is thus sporadic in nature and dependent on the 11-year solar activity cycle, one-half of the 22-year Hale cycle due to the periodic reversal of the Sun's magnetic field (Vainio et al., 2009). For certain considerations, such as the cumulative SEP fluence onto the lunar surface over millennia, time-averaged SEP spectra can be used (e.g., Kuznetsov et al., 2001, or Usoskin and Kovaltsov, 2006), but the total fluence, peak flux, and energy spectra vary greatly between individual SEP events (see comparison by Mewaldt, 2006), and events can last between a few hours and a week (reviewed by Vainio et al., 2009).

The peak flux of GCR particles, at around 500 MeV/nucleon, is 3–4 orders of magnitude lower than the averaged SEP flux at 10 MeV (e.g., Usoskin and Kovaltsov, 2006), but the power law tail of the CR spectra extends up to 10^20 eV at extremely low fluxes. The GCR spectrum is composed of 85% protons, 14% alpha (helium nuclei), and a small fraction of heavy ions (fully ionized atomic nuclei) and electrons. The current paradigm for GCR production is particle acceleration at the expanding shock of supernova remnants (Blandford and Eichler, 1987; Vainio et al., 2009), although details of the source matter and acceleration mechanisms are yet to be perfectly resolved (see the Butt, 2009, study for a recent
review of problems facing the standard model). The source of the ultrahigh-energy CR ($E > 10^{18}$ eV) is unknown and may be extragalactic (see review in Bauleo and Martino, 2009). GCR below about 1 GeV/nucleon are modulated by the heliosphere so their flux is anticorrelated with the solar activity cycle (Klapdor-Kleingrothaus and Zuber, 2000). Thus, GCR and SEP represent two complementary populations of ionizing particles: GCR are present at a relatively constant low flux, but their spectra extend to very high energy levels, whereas SEP have a much higher flux at lower energies and are accelerated within sporadic events.

Figure 1 presents a diagram that summarizes the different CR sources relevant to this discussion: SEP, GCR, trapped particle belts, and induced emissions from CR interactions. A significant radiation source for potentially habitable moons orbiting within the magnetic field of their planet, such as Europa or Titan, are charged particles trapped in circular regions within the dipolar field, which in the terrestrial situation are known as the Van Allen belts (discovered by the Explorer I satellite, Van Allen, 1959, and reviewed recently by Benton and Benton, 2001). Induced emissions are comprised of albedo particles generated by nuclear reactions in CR secondary cascades (explained in more detail below) that backscatter upwards.

Energetic CR primary particles produce extensive showers of secondary particles in Earth’s atmosphere, as shown in Fig. 2. When a GCR strikes an atmospheric nucleus, energetic secondary mesons (pions and kaons), nucleons, gammas, and nuclear fragments are produced, which then interact with other nuclei. Secondary mesons decay over a short timescale to produce muons, gamma rays, and electrons. The air shower is composed of a central “hard component” cone of the electromagnetic cascade (Amsler et al., 2008a). The flux of secondaries builds with increasing shielding depth until it reaches a peak, known as the Pfotzer maximum, after which the average particle energy is below the threshold for new particle production and the cascade is steadily absorbed. In Earth’s atmosphere, the Pfotzer maximum occurs at the altitude of 15–26 km, depending on latitude and solar activity level (Bazilevskaya et al., 2008).

For exposed bodies such as Mars or asteroids, these secondary particle cascades occur in the top meters of subsurface regolith. Nuclear reactions within the secondary cascades generate species such as neutrons and gamma rays that are scattered back out of the subsurface. These backscattered, or albedo, particles can significantly enhance the planetary surface radiation environment (Clowdsey et al., 2001) and can be detected by remote sensing instruments to allow mapping of surface composition. For example, albedo neutrons measured by instruments aboard the Mars Odyssey orbiter provided information on the distribution of hydrogen and, thus, presumably water ice deposits in the martian near surface (Feldman et al., 2002; Mitrofanov et al., 2002).

For more detail on the production and dynamics of SEP and GCR and their interactions with shielding matter (be it planetary atmospheres, surface, or spacecraft pressure hulls), interested readers are directed to recent reviews provided by, for example, Amsler et al. (2008b), Bazilevskaya et al. (2008), and Ferrari and Szuszkiewicz (2009).

3. Biological Effects of Ionizing Radiation

The ionizing radiation field produced by SEP and GCR is harmful to life through direct and indirect mechanisms (Nelson, 2003). Direct damage occurs when deposited energy
excites electrons within biomolecules, which leads to ionization or radiolysis. However, radiation primarily interacts with water, as cells are composed of 40–70% water and bacterial spores 20%, which leads to the indirect mechanism of radiation damage. Irradiated water undergoes ionization and radiolysis, resulting in the production of hydrated free electrons (e\text{aq}/C0) and highly reactive species with unpaired electrons, and free radicals, such as H\textbullet and •OH, or their recombination products such as hydrogen peroxide (Kiefer, 1990). The major reaction pathways of irradiated water are shown in Fig. 3. These radiogenic products are freely diffusible and migrate from their site of production to attack and oxidize biomolecules such as DNA and proteins. If the irradiated water is oxygenated, greater yields of hydrogen peroxide and the hydroxyl radical are produced, and the indirect mechanism of biomolecule damage is enhanced (Baumstark-Khan and Facius, 2001). While ionizing radiation is harmful to life itself, the radiolysis and free-radical chemistry driven by CR on the surface ice of Europa may help maintain an oceanic ecosystem below. This topic is explored in Section 7.9.

Organisms exhibit varying degrees of resilience to ionizing radiation, with the Gram-negative (the classification of bacteria with a kind of cell wall that does not retain crystal violet dye in the Gram staining protocol) nonsporulating bacterium Deinococcus radiodurans being the most radioresistant organism known (Minton, 1994; Battista, 1997; Cox and Battista, 2005). Gamma-irradiated populations of D. radiodurans can survive a dose of 5 kGy without measurable loss of viability despite massive DNA fragmentation, and 1% survival is still found after 10 kGy (Cox and Battista, 2005). The desiccated endospores of bacteria such as Bacillus subtilis are also very radiation resistant (Baltschukat and Horneck, 1991; Setlow, 1994; Moeller et al., 2010), and after 3 kGy of X-rays, 10^{-2} spore viability is still found (Baumstark-Khan and Facius, 2001). For comparison, non-radioreistant bacteria such as E. coli exhibit only of the order 10^{-6} population survival after 1 kGy (Battista, 1997). This remarkable resilience has long been believed to be primarily due to more efficient enzymatic repair of DNA double strand breaks (Battista et al., 1999). More recently, however, evidence has been mounting that other factors such as high intracellular manganese concentrations and protection of proteins from oxidation (Daly et al., 2004; Ghosal et al., 2005; Daly, 2009) are key determinants of D. radiodurans’ radiation survival. Due to this extreme radioresistance, as well as resilience to desiccation, hydrogen peroxide, and UV radiation, D. radiodurans is commonly used as a model strain for microorganisms able to survive on the martian surface (e.g., Richmond et al., 1999; Horneck, 2000; Pavlov et al., 2002; Dartnell et al., 2007a; de La Vega et al., 2007). More recently, strains of a hyperthermophilic archaeon Thermococcus, isolated from a submarine hydrothermal vent environment, have been found with a radiation resistance approaching that of D. radiodurans (Jolivet et al., 2003, 2004).

The effects of ionizing radiation on the structure of biomolecules and the survival of organisms has been well covered elsewhere, so only a brief summary has been provided here to offer context for the following discussion on astrobiological aspects of ionizing radiation that have not yet received due coverage. For more complete reviews of space radiobiology, the interested reader is referred to reviews by Baumstark-Khan and Facius (2001) or Nelson (2003) and

![FIG. 2. The extensive cascade of secondary particles, or air shower, generated by an energetic primary CR striking a nucleus in Earth’s upper atmosphere. The central hadronic core is composed of neutrons, protons, and nuclear fragments propagating essentially straight ahead, within the spreading electromagnetic cone of pions and muons and their decay products electrons and gamma-ray photons. Particle flux builds to a peak, known as the Pfotzer maximum, before being increasingly absorbed. Color images available online at www.liebertonline.com/ast](image-url)
papers that consider more specifically the hazard posed by CR to human space exploration (Todd, 2003; NCRP, 2006).

Ionizing radiation can induce mutations and act as a selective filter for survival, and so it can potentially influence both of these drivers of natural selection and evolution of life (Committee on the Origins and Evolution of Life, 2005). Except for transient extreme ionization events (see Section 8), Earth’s surface has been well shielded from CR throughout its history by a thick atmosphere (Svensmark, 2006, estimated variation in CR flux over the planet’s history; discussed here in Section 8.8), and although organism exposure from environmental radionuclides has decreased by a factor of around 5 since primordial Earth, this has largely been counteracted by rising oxygen levels and consequent enhancement of radiation damage (Karam and Leslie, 1999; Karam et al., 2001). Surface radiation contributes to mutagenesis rates along with a variety of other causes, including replication errors and DNA damage from UV, chemical oxidants, and desiccation (Makarova et al., 2001). The rate of DNA damage wreaked by reactive oxygen species (such as hydroxyl radicals or hydrogen peroxide) generated by an aerobic cell’s own metabolism is, however, estimated to be orders of magnitude greater than that due to current background levels of ionizing radiation (Imlay, 2003).

While this may be true for background levels of ionizing radiation, stochastic extreme cosmic ionization events, such as a nearby supernova, could occasionally and transiently raise the mutation rate of exposed organisms. The question is whether such pulses of hypermutation would be generally deleterious to a population, potentially resulting even in extinction, or could perhaps prove beneficial and enhance evolution. Smith et al. (2004a) argued that a fluctuating radiation environment and hypermutation events would produce a more diverse adaptive landscape and offer natural selection greater allelic variation within a population to operate upon. They thus concluded that habitable worlds with a more strongly fluctuating radiation environment than Earth may in fact experience accelerated evolution and development of complexity. Smith et al. (2004a) presented an in-depth account of the argument, as did Scalo et al. (2007, and references therein), who considered the possibility of accelerated evolution on planets around M-class stars driven by a fluctuating radiation environment. There is genomic evidence from bacteria residing inside eukaryotic cells (endocellular symbionts) that increased mutation rates can indeed drive faster evolution (Itoh et al., 2002). Organisms with mutator phenotypes (exhibiting elevated mutation rates) are known in nature, and though advantageous in the short term when colonizing or adapting to a new environment, they are generally maladaptive due to the accumulation of deleterious mutations (e.g., Giraud et al., 2001). The concept of adaptively variable mutation rates and the ability to evolve, or “evolvability,” of organisms was reviewed by Radman et al. (1999) and Pigliucci (2008). However, there is no experimental evidence that explicitly links heightened radiation exposure and subsequent hypermutation to an enhanced rate of evolution. Indeed, bacteria chronically exposed to high doses of ionizing radiation were found to adapt to exhibit greater radiation resistance and thus maintain a lower mutation rate (Ewing, 1995, 1997).

While background CR currently pose no significant hazard to the survival of life on Earth’s surface, cosmic radiation may well play a decisive role in the terrestrial biosphere during acute astrophysical events (see Section 8), or over the longer term on planets with less shielding, such as Mars (see Section 7.8), or habitable planets orbiting other classes of stars (see Section 7.2). There is also the possibility that a high flux of ionizing radiation could support an extraterrestrial ecology on radiogenic products (see Section 7.9). These biological ramifications of ionizing radiation will be treated later, but for now this review will discuss the ways in which CR influence the creation of habitable environments for life, from their role in the formation of stars to the synthesis of biogenic elements and the organic precursor molecules of biochemistry.
4. The Role of Cosmic Rays in Star Formation

Cosmic rays pervade the Galaxy and ionize the atoms and molecules of the interstellar medium (and are the dominant source of ionization in deep regions where UV radiation cannot penetrate to photoionize; McKee and Ostriker, 2007; Padovani et al., 2009). When charged, these interstellar gas particles are influenced by galactic magnetic fields and through ion-neutral collisions with other particles affect the dynamics of the entire gas cloud. So GCR help tie the behavior of interstellar gas to the ambient magnetic fields of the Galaxy. Magnetic fields stabilize interstellar molecular clouds against gravitational collapse, which regulates the collapse of protostellar clumps (Zweibel and Heiles, 1997) and sheds angular momentum of star-forming cloud cores and protostellar disks by magnetic braking (Desch and Mouschovias, 2001; Allen et al., 2003; Larson, 2003). The ramification of this is that, alongside UV photoionization, ionization by CR plays a role in star formation (reviewed recently by McKee and Ostriker, 2007; Larson, 2010) and is particularly significant deep in star-forming clouds or in the outer regions of protostellar disks. Cosmic rays also act to heat the interstellar medium (ISM), which raises the gas pressure and further increases the stability against gravitational collapse (Zweibel and Heiles, 1997; Padovani et al., 2009). This coupling of interstellar gas to galactic magnetic fields by CR ionization is therefore thought to exert control over the mass distribution of star formation (Rees, 1987).

Thus, it could be argued that each step of the cycle of stellar evolution in the Galaxy is linked together by CR, as illustrated in Fig. 4. Supernovae shock fronts accelerate the particles of CR, which are contained by the galactic magnetic field. These trapped CR act to ionize the interstellar gas clouds so that they become responsive to the magnetic fields. This process moderates star formation, with massive stars ending their evolution in supernovae, thus completing the cycle.

Creation of the first stars, termed Population III stars, in the primordial Milky Way and other galaxies would have necessitated different mechanisms from modern star birth. Such early stars must have formed from metal-free primordial gas in the absence of dust grains (and thus without efficient radiative cooling promoting contraction) and without ionization of the gas by CR or regulation by galactic magnetic fields. Modeling studies indicate that the earliest stars in the Universe were thus probably very massive; in excess of 100 times solar mass (Abel et al., 2002; Bromm et al., 2002; Tan and McKee, 2004, 2008; and recently reviewed by Bromm et al., 2009). Observational evidence for these early massive stars is anticipated over the coming decade (Bromm et al., 2009), aided by the James Webb Space Telescope. The luminous supernova SN 2007bi is believed to have been a pair-instability supernova, predicted by theory to result only from stars greater than 140 solar mass, which suggests the formation of very massive stars in modern dwarf galaxies through similar mechanisms to the first stars in the Universe (Gal-Yam et al., 2009).

After the first generation of supernovae, interstellar space would have become pervaded with GCR and enriched with heavier elements. Not all the heavy elements of the Periodic Table are produced by stellar or supernova nucleosynthesis, however, and the formation of several key elements can only be explained by CR, as discussed in the next section.

5. Role of Cosmic Rays in Biogenic Elements

The lightest isotopes, H, H-2, He-3, He-4, and a small amount of Li-7, were produced by nucleosynthesis during the Big Bang. The heavier elements are produced by nucleosynthesis reactions in stellar fusion, such as the triple alpha process, which fuses three helium nuclei into carbon; and elements beyond iron are generated by supernova nucleosynthesis. The cosmic abundance of the light elements lithium, beryllium, and boron (LiBeB) is, however, initially difficult to explain as they are destroyed by the high-temperature environments of the early Universe and stellar interiors (Prantzos et al., 1993). These LiBeB elements play important roles in biology, so the mechanism of their generation is important here.

A clue as to the origin of these fragile, light elements is provided by the observation that the abundance of this triplet relative to hydrogen is many orders of magnitude greater in GCR compared to the Solar System composition, as shown in Fig. 5.

The source of these fragile LiBeB elements was proposed 40 years ago (Reeves et al., 1970; Meneguzzi et al., 1971) to be nuclear spallation reactions between the energetic protons and alpha particles of the GCR and carbon, nitrogen, and oxygen nuclei of the low-density ISM. Another source was suggested to be fast carbon and oxygen nuclei GCR particles that impact the ambient hydrogen and helium (Reeves, 1994; Zhai, 1995; Vangioni-Flam and Casse, 2000). One particular problem in fitting such models to observed isotopic ratios is the fact that the metallicity of the ISM has increased over galactic history (through stellar and supernova nucleosynthesis) and modeled GCR LiBeB synthesis calculations either underproduce Be and B at low metallicities or overproduce Li (mostly produced by He+He reactions) (Prantzos et al.,...
1993). Proposed explanations for the inconsistencies between predicted and observed ratios of LiBeB isotopes include refinements to the sources of CR (Vangioni-Flam and Casse, 2000) or additional sources of LiBeB isotopes, such as the hot outer envelopes of asymptotic giant branch stars (“red giants;” Prantzos, 2006). Ramaty et al. (2001) and Prantzos (2007) presented comprehensive reviews of the development of theories to explain GCR production of the LiBeB elements during the evolution of the Galaxy.

These CR-generated LiBeB elements, incorporated into new stars and planets from collapsing interstellar clouds, play important roles within biology. The elements employed by terrestrial life are, on the whole, those that are most abundant in Earth’s crust, ocean, and atmosphere. Lithium and beryllium are less abundant at Earth’s surface than alternative elements such as sodium, potassium, magnesium, and calcium, all of which are able to perform the same chemical functions (Fraústo da Silva and Williams, 2001). So lithium and beryllium are not essential elements for terrestrial life and can in fact show toxicity at high concentrations. They form small ions similar to Mg$^{2+}$, and so they can interfere with biochemical processes through their interaction with phosphate-containing molecules (Fraústo da Silva and Williams, 2001). Beryllium has also been implicated in enzyme inactivation and malfunction (Wackett et al., 2004).

Boron, the third element of the CR-produced LiBeB triplet, is, however, vital for the proper biochemical functionality of many diverse branches of terrestrial life (reviewed by Bolaños et al., 2004). For example, boron is required by cyanobacteria for nitrogen fixation, plants employ borate esters in keeping their tissues flexible (Wackett et al., 2004), and boron is also considered a necessary ultratrace element in human metabolism (Nielsen, 1998). The main biochemical functionality of boron is believed to be its ability to stabilize molecules with cis-diol groups (where both –OH functional groups are on the same side; Bolaños et al., 2004), such as cyclic ribose (Ricardo et al., 2004), and it is for this ability (Benner et al., 2010) that boron has also been hypothesized to have played a crucial role in prebiotic chemistry and the origin of life. The conditions for an RNA world that led to cellular life have been doubted, as the prebiotic synthesis of ribose and other pentose sugars is difficult to explain. The presence of borate minerals, however, stabilizes the 1,2-diol structure of these pentoses and allows them to accumulate under prebiotic conditions (Ricardo et al., 2004; Benner et al., 2010). Grew et al. (2011) discussed the availability of borate minerals around 4 billion years ago to support an RNA world, and concluded that it critically depends on an early differentiation of a granitic continental crust.

Schulze-Makuch and Irwin (2008) even discussed the possibility of boron forming the basis of an extraterrestrial biochemistry. Boron, like carbon, has a remarkable propensity for forming covalent molecular compounds, and boron-nitrogen compounds exhibit physical and chemical properties similar to alkanes and aromatic hydrocarbons. Boron compounds dissolve well in nitrogen and ammonia, so these could serve as the biosolvent for low-temperature boron-based life (Schulze-Makuch and Irwin, 2008).

So CR are crucial for the creation of the light LiBeB elements that are all biologically active, and boron is vital for biochemistry and may have played a key role in the origin of life. As well as driving nuclear reactions to generate elements crucial for biology, CR play a major role in the interstellar synthesis of organic molecules for life, as will be seen in the next section.
6. Role of Cosmic Rays in Astrochemistry and Prebiotic Organic Molecules

Despite the low densities and temperatures, a great deal of complex chemistry occurs within interstellar dense molecular clouds. Cosmic rays, as well as UV photoionization, drive these reactions, and many of the molecules created here are crucial for biology (for example, see recent reviews by Ehrenfreund and Charnley, 2000; Charnley et al., 2002; and Snow and Bierbaum, 2008). Much of this “astrochemistry” is thought to depend on surface catalysis on interstellar dust grains (as reviewed by Williams and Herbst, 2002). Reactants and products are in a dynamic equilibrium of adsorption onto grain surfaces and release back into the gas phase through desorption induced by UV and CR (Öberg et al., 2007). Heating of these dark molecular clouds, and thus the thermal energy required for surface jumping and chemistry, is provided by CR. Cosmic ray ionization also drives the ion-molecule chemical reactions for atoms whose ionization potential is greater than that of hydrogen and, in particular, initiates the reactions of nitrogen- and oxygen-containing molecules (Federman et al., 1996). Ultraviolet photoionization also plays a role in driving astrochemistry in the diffuse ISM. Within cold dark molecular clouds (~10 K), however, the dense dust absorbs UV photons, and only CR can penetrate inside (reviewed well by Dalgarno, 2006). Cosmic ray excitation of molecular hydrogen within the clouds also produces UV fluorescence (Öberg et al., 2007), and this re-emitted energy contribution is thought to be roughly equal to that of direct CR ionization in driving astrochemistry (Moore et al., 2001).

Molecular species such as water, carbon dioxide, methanol, and formaldehyde are produced and form an icy mantle of volatiles around the grains where further irradiation processing occurs. Once mantles have built up, chemistry is not restricted to the surface as reactants can diffuse vertically through the ice layers (Williams and Herbst, 2002). Simulation and laboratory work on irradiation of ices have demonstrated that relatively simple ices can be processed into much more complex molecules such as ethanol, carbonic acid, lactic acid, and long carbon chains (Charnley et al., 2002; Shaw, 2007). Simakov et al. (2002) also showed that nucleotides (the monomer components of DNA and RNA) can be produced by high-energy proton irradiation of simulated interstellar dust grains that contain nucleosides and inorganic phosphate. Although laboratory experiments of ion-irradiated ices have produced many complex organic molecules, the exact nature of the CR flux present in different regions of molecular clouds is not well understood (Moore et al., 2001; Indriolo et al., 2007). Water ice trapped in lunar polar craters and exposed to the CR flux throughout the history of the Solar System offers the possibility of field-testing models of astrochemistry and radiation-driven synthesis of organics on icy planetary surfaces and dust grains in interstellar clouds (Lucey, 2000).

As cores within interstellar dust clouds collapse to start the process of star formation, a revolving accretion disk of gas and dust forms around the central protostar. Such a protoplanetary disk is much denser and warmer than the surrounding molecular cloud, and further chemical processing occurs during this stage (see van Dishoeck and Blake, 2003). Organic molecules produced in this circumstellar region become incorporated into comets and meteorites, forming a pool of prebiotic compounds that can be delivered to habitable planets and moons and hence help initiate an origin of life (Ehrenfreund and Charnley, 2000).

Molecular species in interstellar and circumstellar regions have been identified by both their rotational transitions (radiofrequency spectroscopy) and vibrational transitions (IR spectroscopy), and to date over 140 different species have been confirmed (Snow and Bierbaum, 2008). Prebiotic molecules such as amino acids and sugars are expected from theoretical modeling and laboratory experiments but are difficult to detect due to their low abundance in the gas phase (Herbst and van Dishoeck, 2009, presented an excellent review of complex organic interstellar molecules). The simplest sugar, glycolaldehyde, has been independently detected both toward the galactic core (Hollis et al., 2000) and in hot star-forming regions (Beltrán et al., 2009). However, the detection of glycine, the simplest amino acid, was claimed by Kuan et al. (2003) in several hot star-forming regions, but this interpretation has since been contested by the spectral analysis of Snyder et al. (2005). Nevertheless, the increasing realization over recent decades has been the enormous degree of interstellar chemistry, driven overwhelmingly by CR ionization deep inside molecular clouds in addition to UV photochemistry at later stages, which produces molecules of great importance to biology.

Taken to much more speculative lengths, it has been conjectured that, in addition to ionization-driven organic chemistry within irradiated ices, CR may even allow simple living systems to survive in a solid state. Frozen within cometary ice, such organisms may occasionally derive metabolic energy from the tracks of free radicals created by the propagation of CR particles, metabolizing very slowly at perhaps only a few transformations per millennium (Committee on the Limits of Organic Life in Planetary Systems, 2007).

Cosmic ray ionization and driving of chemistry remains a very important process for planetary atmospheres and the icy surface of comets or moons like Europa, as will be addressed later. For now, I will discuss the dynamic nature and shielding of the flux of CR onto habitable planets and moons.

7. Influence of Cosmic Rays on Planetary Habitability

7.1. Atmosphere

During collapse of the young stellar object and accretion of planets, asteroids, and comets from the protoplanetary disk, ionization from CR continues to play a key role in driving astrochemistry in the forming solar system. Once the central star has formed, however, it generates a magnetic field and a fast outflowing stream of charged particles, the stellar wind, which blows a cavity in the ISM. The Sun’s heliosphere (or astrosphere for the general case) moderates the entry of GCR into the Solar System as a function of particle energy and charge, and also in a time-varying way due to the 11-year cycles of solar activity and “grand minima” of activity such as the Maunder Minimum (1645–1715). Transient and local reductions, known as Forbush decreases, are caused by the shielding effect of the magnetic field carried by coronal mass ejections (reviewed in Cane, 2000). Scherer et al. (2006) and Vainio et al. (2009) presented good reviews of the variability of the particle radiation environment at Earth. Our
understanding is rather limited, and model dependent, of what the unmodified spectra of CR in interstellar space actually are, as in situ particle energy measurements are only now becoming available as first Voyager 1 on 16 December 2004 (Stone et al., 2005) and then Voyager 2 on 30 August 2007 (Stone et al., 2008) passed the termination shock of the heliosphere (the position of which fluctuates with the strength of the solar wind; Webber and Intriligator, 2011).

The extent of this atmosphere and thus how much protection from GCR is afforded to habitable planets is determined by the balance between the outward pressure of the stellar wind and the density of the local ISM the star is passing through (for good overviews see studies by Scherer et al., 2002, and Fahr, 2004). Passage of the Solar System through dense interstellar clouds would compress the heliosphere and significantly change the interplanetary environment Earth is exposed to (Scherer et al., 2002). This may have occurred in very recent geological history. Zank and Frisch (1999) modeled the interaction between the heliosphere and interstellar clouds and argued that anomalous peaks in Be-10 concentration discovered in Antarctic ice cores that correspond to 33,000 and 60,000 years ago are caused by enhancements in CR flux arriving at Earth’s atmosphere due to reduced heliospheric shielding.

It is predicted that dense interstellar clouds could even compress the heliosphere to less than 1 AU, which would expose Earth to the unmodulated flux of GCR as well as infalling interstellar gas and dust. Such an extreme phenomenon has been termed “astrospheric collapse” or “descreening” (Smith and Scalo, 2009). Biological hazards from the loss of heliospheric protection include increased surface UV flux due to stratospheric ozone destruction from increased CR flux (Pavlov et al., 2005a) or interstellar hydrogen gas flowing into the atmosphere (Yeghikyan and Fahr, 2004a, 2004b), and interstellar dust triggering runaway glaciation (Pavlov et al., 2005b). Such deleterious outcomes are expected to far outweigh any likely benefits of diminished SCR flux onto Earth during astrospheric collapse.

Smith and Scalo (2009) explored the parameter space of both ISM density and stellar mass to calculate how often planets in the habitable zone of their star would experience such astrosphere descreening events and so be vulnerable to the biological hazards listed above. They found that the risk of descreening decreases strongly with decreasing stellar mass, due both to the weaker focusing of the interstellar flow by the star’s gravity and the fact that the habitable zone is closer for lower-mass stars. Descreening could befall a habitable planet around a Sun-type star up to 10 times per billion years, but those orbiting M-class dwarves are virtually never exposed to the severe consequences of astrospheric collapse (Smith and Scalo, 2009).

The shockwave from a nearby supernova explosion could also potentially overcome the solar wind ram pressure and compress the heliosphere to less than 1 AU. Fields et al. (2008) performed numerical hydrodynamical modeling and calculated that a supernova 20 parsecs (pc) distant can penetrate the Solar System to within a few astronomical units, and that material from an event 10 pc distant can collapse the heliosphere to almost exactly 1 AU. The debris from a recent nearby supernova, identified by the isotopic signature of radioactive $^{56}$Fe, has been found in ~2-3 million-year-old marine sediments (Knief et al., 1999, 2004), and the event may have caused such a significant compression of the heliosphere (Fields et al., 2008).

### 7.2. Planetary magnetosphere

The astrosphere is the first line of shielding against GCR in a planetary system. Many planets, and even moons (e.g., Ganymede; Kivelson et al., 1996), generate a global dipolar magnetic field, and this planetary magnetosphere further screens the spectrum of CR able to reach the atmosphere or surface.

The rigidity of a charged particle is defined as the ratio between its relativistic momentum and charge, and particles with a higher rigidity are deflected to a lesser extent by a magnetic field. The transmission of CR through a magnetosphere can be quantified by the cut-off rigidity, which is the minimum rigidity particles require to reach a particular position within the magnetosphere from a particular arrival direction (Vainio et al., 2009). The magnetic equator of a planetary surface protected by a global dipolar magnetic field has the highest cut-off rigidity, whereas particles with a much lower minimum rigidity are able to access the polar regions. Unmagnetized bodies present no such obstacle, and the CR flux is independent of latitude.

Grießmeier et al. (2005) analyzed the magnetospheric protection an Earth-like planet is likely to receive against CR within the habitable zone of a low-mass star, 0.5 $M_{\odot}$ (on the border between M class and K class). Orbiting at only 0.2 AU from its star, such a terrestrial planet would become tidally locked, which would result in a rotation rate 2% that of Earth and consequently a reduced magnetic moment of the global dipole field. The effects on the magnetosphere of such a planetary system are twofold. First, such a slowly rotating planet will generate a weaker magnetic field than Earth (which also has implications for protection from atmospheric erosion by the solar wind; Khodachenko et al., 2007). Second, the habitable zone around a star less massive than the Sun is much closer, so the Earth-like planet is exposed to a more intense stellar wind ram pressure and thus compression of the planetary magnetosphere. Grießmeier et al. (2005) therefore calculated that surface life on a habitable Earth-like planet around an M-class star would be exposed to a higher flux of GCR, which would also have access to a greater area of the planetary surface. Extending this analysis, Grießmeier et al. (2009) showed that the variation in stellar wind pressure with orbital distance has little influence on the shielding of GCR and that the first effect, tidal locking and reduced magnetic moment, is in fact dominant. They also found that the composition and average density of the planet affects the modeled magnetic moment, with a rocky super-Earth receiving 2–3 times more protection against CR >0.0 MeV than an ocean planet of the same mass.

One effect of a reduced magnetic moment and greater GCR penetration may be more widespread loss of the atmospheric ozone shield, as can only occur in polar regions on Earth where the magnetosphere allows energetic particles into the upper atmosphere (Scalo et al., 2007), although in their modeling work Grenfell et al. (2007) found that GCR-driven atmospheric chemistry only minimally affects ozone and water concentrations (also see Section 7.11 of this paper). Life may face no direct radiological hazard from enhanced GCR flux, as Grießmeier et al. (2005) pointed out that a dense
atmosphere would absorb CR regardless of magnetospheric deflection, a result that has also been discussed for the case of primordial Mars (Molina-Cuberos et al., 2001; Dartnell et al., 2007b; Schneider and Kasting, 2009).

The local source of CR—energetic particles accelerated by the active star—is, however, likely to be a significant threat to habitable planets orbiting young low-mass stars. For the first billion years or so, M-class stars are very active and produce frequent and intense flares and energetic particle events (Scalo et al., 2007), which could produce a SCR flux orders of magnitude higher than that incident upon Earth (Lammer et al., 2009). Habitable M-star planets would thus be exposed to recurrent pulses of SCR, with likely implications for life in the radiation environment created by energetic particles penetrating to the surface (Grießmeier et al., 2005), as well as erosion of the atmosphere (Khotidenko et al., 2007).

Just as the shielding of GCR provided by the atmosphere is time variable, over timescales of both the activity cycle of stars as well as passage through dense interstellar clouds with galactic orbital motion, the shielding provided by a planetary magnetosphere is not constant. Earth's magnetic field reversals can temporarily affect atmospheric chemistry through increased CR penetration (reviewed by Glassmeier and Vogt, 2010), especially if this occurs coincident with another phenomenon such as a nearby supernova, although that is improbable (Crutzen and Brühl, 1996). The magnetic field intensity also fluctuates over geological timescales. Yamazaki and Oda (2002) found a 100,000-year periodicity in the inclination and intensity of the dipolar field over the past 2.25 million years, with the field intensity peaking at no more than twice that of today. This variability is believed to be modulated by the eccentricity of Earth's orbit and may affect CR penetration into the atmosphere.

Looking at longer timescales, except for transient periods of reversal, the geomagnetic field strength is thought to have varied in intensity by approximately \( \pm 50\% \) from its current magnetic moment of around \( 8 \times 10^{22} \text{Am}^2 \) over the past 160 million years (paleomagnetic records and models were reviewed by Aubert et al., 2010). Tarduno et al. (2010) reported results on the Archean geodynamo field strength recorded by magnetic inclusions in single silicate crystals from the Kaapvaal craton in South Africa. They found that 3.4–3.45 billion years ago the geomagnetic field was around 50–70\% that of the present day. This slightly weaker field in the Archean, combined with a greater solar wind ram pressure from a younger, more active Sun, suggests a more compressed magnetosphere. Tarduno et al. (2010) calculated that the magnetopause stand-off distance at that time would have been less than about five Earth radii, which is similar to the effect of a large coronal mass ejection today.

7.3. Interaction of cosmic rays with atmosphere

The atmosphere of a planetary body (including large satellites such as Titan) presents the final level of shielding of the surface from ionizing CR with sufficient energy to overcome both the atmosphere and planetary magnetosphere. Although the energy delivered to Earth's atmosphere by GCR is tiny compared to solar electromagnetic radiation, which is greater by a factor of \( 10^6 \), they are the only source of ionization at altitudes between 3 and 35 km (Stozhkov, 2003) and so are a crucial driver of atmospheric chemistry and meteorology. Absorption of this energy flux results in a vertical profile of excitation and ionization of atomic and molecular species within the atmosphere.

An ionosphere is thus created around a planet or moon by the ionization of its neutral atmosphere. In many situations, photoionization by solar UV photons is the dominant source of atmospheric ions, but energetic particle radiation also plays a major role (notably for the nightside ionosphere and bodies at a greater distance from the Sun). Ionization can be caused by charged particles trapped in a planet's dipolar magnetic field precipitating into the upper atmosphere around the poles (e.g., Earth; reviewed by Lyons, 1997), or a moon if it is embedded in its parent planetary magnetosphere (e.g., Titan), as well as GCR and SEP cosmic rays arriving from beyond the magnetosphere (Kivelson and Russell, 1995; also see a recent review of solar system ionospheres by Witasse et al., 2008). Atmospheric ionization is very dynamic and variable over both space and time, being influenced in the short term and longer term by solar UV output and flares; CR variability; gamma-ray, X-ray, and energetic particle influx from supernovae in the stellar neighborhood; and passage of the Solar System through dense molecular clouds (Vasyliov et al., 2008).

Ionization induced by CR penetrates much deeper into the atmosphere than that from solar UV photoionization; so planetary bodies such as Earth, Venus, and Titan exhibit a second ionosphere below that created by solar UV, which corresponds to the Pfitzer maximum of secondary cascade intensity. For example, the upper ionosphere of Titan, created by solar UV and Saturn's trapped particle belt, has a daytime peak electron density of 3000–6000 cm\(^{-3}\) (depending on solar zenith angle and impact of magnetospheric electrons) at an altitude of around 1000 km, as indicated by modeling studies (Ip, 1990; Roboz and Nagy, 1994; Keller et al., 1998; Banaszkiewicz et al., 2000; Cravens et al., 2004; Wilson and Atreya, 2004; Krasnopolsky, 2009) and corroborated by Voyager 1 radio occultation observations (Bird et al., 1997) and Cassini measurements (Wahlund et al., 2005). The lower, GCR-generated ionosphere has a slightly lower peak electron density of 1000–2000 cm\(^{-3}\) at an altitude of 70–90 km (modeled by Capone et al., 1980; Borucki et al., 1987, 2006; Molina-Cuberos et al., 1999a, 1999b).

This lower ionization peak from CR is the dominant source of ions in the terrestrial troposphere and drives atmospheric chemistry and potentially weather and climate dynamics. Several recent papers have modeled the ionization profile through Earth's atmosphere from SCR and GCR (for example, see Jackman et al., 2000; Usoskin and Kovalsov, 2006; Vasilyev et al., 2008), with up-to-date overviews of the field by Bazilevskaya et al. (2008) and Usoskin et al. (2009), these authors comparing the results of several different numerical models.

As well as a potential role of atmospheric cascades in triggering lightning on Earth (Gurevich et al., 1999; Dwyer, 2005), CR ionization has been implicated in driving variability in global climate and the production of organic molecules in the primordial atmosphere, which will be explored in the next two subsections.

7.4. Cosmic rays and climate forcing

There exists a long-running debate on whether variations in the CR flux incident on Earth can affect global climate,
principally by driving cloud formation. The effect of cloud cover is complicated by the fact that it depends on the type; an increase of low-altitude optically thick clouds will result in a cooling through greater albedo and reflection of solar radiation, whereas an increase of high altitude optically thin clouds will warm the planet by trapping reradiated IR. The claimed link between CR flux and climate is that increased ionization of the troposphere by a greater flux of GCR entering the inner Solar System, such as during a period of reduced solar activity, generates additional nucleation sites for water condensation. This increased abundance of nucleation sites causes an increase in the formation of low-altitude clouds and so drives global cooling. For reviews of this hypothesis, see studies by Svensmark (2000, 2007), Carslaw et al. (2002), and Kirkby (2007). However, the actual physical and chemical mechanisms by which CR can influence condensation nuclei and thus cloud cover are themselves unclear (as reviewed by Carslaw et al., 2002; Kirkby, 2007; Bailer-Jones, 2009). So while an observed correlation between CR intensity and average cloud cover over Earth during the last solar cycle has been reported, critics either question the validity of the correlation (for example, Laut, 2003 presented a cautionary analysis of several claimed correlations between solar activity and climate change) or accept that the correlation is real but explain it by other periodic influences (reviewed by Carslaw et al., 2002). For example, in their simulations of the “ion-aerosol clear-air” hypothesis, Pierce and Adams (2009) found that changes in cloud condensation nuclei from variation in CR over the course of the 11-year solar cycle are 2 orders of magnitude too small to account for the observed changes in low-level cloud cover, but they said that a connection between CR and clouds could operate by other mechanisms. However, looking at shorter timescales, Svensmark et al. (2009) reported a clear link between CR variability during Forbush decreases and measurements of both atmospheric aerosol abundance and low clouds. For further discussion on the potential link between the solar activity–modulated CR flux at Earth, tropospheric ionization, cloud formation, and climate see studies by Svensmark and Friis-Christensen (1997), Bazilevskaya et al. (2000), Carslaw et al. (2002), Shaviv (2005), Section 15 of Scherer et al. (2006), Kirkby (2007), and Bailer-Jones (2009).

7.5. Cosmic ray production of organics within the atmosphere

It is clear that excitation and ionization from CR has a major influence on pushing the atmospheric chemistry of planetary bodies. This energy source generates complex organic species in interstellar molecular clouds, and it is also expected to produce biologically relevant precursor molecules in situ on habitable worlds, including primordial Earth and the complex carbon-nitrogen chemistry and tholin production in Titan’s atmosphere (discussed in Section 7.7).

Kobayashi and Tsuji (1997) used high-energy protons to irradiate a gas mixture of carbon monoxide, nitrogen, and water, which is believed to be characteristic of the primordial terrestrial atmosphere, and observed the production of uracil, one of the bases of RNA. In subsequent experiments (Kobayashi et al., 1998, 1999), a similar atmospheric composition was irradiated with high-energy electrons, protons, and helium nuclei, which produced a wide range of amino acids.

Beyond Earth, atmospheric ionization processes are important for other potentially habitable planets and moons in our Solar System, as well as exoplanets and exomoons. Only two potential astrobiological habitats in the Solar System have substantial atmospheres, Venus and Titan, which are discussed next. Studies in which the various effects of extreme atmospheric ionization on the terrestrial environment were analyzed, including ozone destruction and climate forcing, are treated in Section 8.

7.6. Venus atmosphere ionization

Venus possesses an atmosphere more dense than that of Earth, but due to the lack of a global magnetic field, CR particles can penetrate further and create an ionization profile deeper into the venusian atmosphere (Upadhyay et al., 1994). Ionizing solar UV does not penetrate much deeper into the venusian atmosphere than ~120 km altitude, particle radiation from solar flares dominate in the 60–100 km altitude range, and below 60 km GCR are the primary source of ionization (Borucki et al., 1982). This CR-induced ionosphere has been modeled by Dubach et al. (1974), Borucki et al. (1982), Upadhyay et al. (1994), Upadhyay and Singh (1995), and Martin-Torres and Molina-Cuberos (2002).

The total shielding thickness of the venusian atmosphere, around 10^5 g/cm^2 (Borucki et al., 1982), is two orders of magnitude greater than the terrestrial shield of 1033 g/cm^2 (United States Committee on Extension to the Standard Atmosphere NOAA, 1976), so even the most energetic CR
present no biological hazard on the surface. The environmental regime of the current venusian surface, however, is not suitable for organic life, and it is the lower or middle global cloud deck that has been proposed as a potential atmospheric habitable zone (e.g., Cockell, 1999; Schulze-Makuch et al., 2004). Between 48 and 57 km altitude, local physical conditions are conducive to life, with atmospheric pressure between 0.1 and 1 bar and temperature within the range 0–60°C. In this region, water is available in long-lasting cloud droplets—although of very low pH from high concentrations of sulfuric acid—and the remaining solar UV (less than a few percent of the incident flux penetrates to the lower cloud level; Cockell, 1999) may be effectively screened or, perhaps, photosynthetically harvested by microorganisms encased in S8 (Schulze-Makuch et al., 2004).

Within this putative aerial biosphere, a further survival hazard which may prove to be significant is the radiation environment generated by CR. The peak in CR ionization occurs at an altitude of 25–65 km in the venusian atmosphere (Upadhyay et al., 1994; Upadhyay and Singh, 1995), which overlaps with the potential habitable zone, as shown in Fig. 6.

The peak particle intensity in Earth’s atmosphere, the Pfotzer maximum, occurs at 15–26 km altitude, depending on latitude and solar activity level (Bazilevskaya et al., 2008). The Concorde used to fly at a cruise altitude of 18 km, well within this radiation peak, and civil aviation still operates typically at 10–12 km altitude. The total radiation dose from GCR received by a member of the aircrew working a reasonable workload of 700 flight hours per year on subsonic aircraft, including polar routes, is considerably lower than the occupational limit recommended for nonpregnant adults by the International Commission on Radiological Protection (ICRP, 1991) of 20 mSv/year averaged over 5 years (O’Brien et al., 1996). However, the more conservative recommendation for a pregnant woman of no more than 2 mSv over the pregnancy (ICRP, 1991), and not greater than 0.5 mSv in any month (NCRP, 1993), can be exceeded during a standard flight schedule on high-latitude routes (O’Brien et al., 1996; Shea and Smart, 2000). Passengers and crew aboard Concorde were exposed to an equivalent dose rate from GCR 5–12 times higher than subsonic, lower-altitude, short-haul flights (reviewed recently in Bagshaw, 2008). Every Concorde was fitted with a radiation monitor, and the pilot was alerted to descend to a lower altitude if the dose rate exceeded safety limits during a solar particle event (for details, see Lim, 2002). It is possible that the radiation environment in the Pfotzer maximum of the venusian atmosphere may also represent a survival threat to life already stressed by other environmental hazards such as acidity and solar UV. The radiation hazard would be particularly acute during a large solar flare, especially considering Venus orbits 30% closer to the Sun than Earth. This potential radiation stress to a venusian aerial ecosystem has not been fully considered before and may prove significant with further study.

7.7. Titan atmosphere ionization

The dense atmosphere of Titan (total shielding depth of 11,000 g/cm²; integral of standard atmosphere density

FIG. 6. Schematic of the venusian atmospheric profile, which displays the upper and lower bounds of the three venusian cloud decks and the extent of the aerial habitable zone (green), defined by the region between 0°C and 150°C (dashed lines projected from red temperature plot). A venusian biosphere could reside in the overlap region between the cloud layer and habitable temperature regime but lies near the peak in atmospheric ionization from cosmic radiation (blue). Diagram constructed using temperature profile from Prinn and Fegley (1987), CR ionization modeling from Upadhyay et al. (1994), and habitable zone discussion in Cockell (1999). Color images available online at www.liebertonline.com/ast
profile in Yelle et al., 1997, and stated as such by Krasnopolsky, 2009) is believed to have a similar chemistry to Earth’s primordial atmosphere (Molina-Cuberos et al., 1999b). This atmospheric chemistry is driven at high altitudes by ionization from solar UV and trapped particles in the magnetosphere of Saturn, creating an ionosphere at an altitude above about 800 km (Molina-Cuberos et al., 1999a; Cravens et al., 2006), the composition of which has now been measured (Cravens et al., 2006; Coates et al., 2007) by the Cassini orbiter (reviewed recently by Brown et al., 2009). Solar UV irradiance in the outer Solar System is much lower than in the solar wind, and Titan’s atmosphere, with its composition of nitrogen, methane, and hydrocarbons, is very opaque to UV. Saturn’s magnetic field at Titan is not powerful enough to shield CR flux significantly, and Titan generates no intrinsic magnetic field of its own (Backes et al., 2005), so the Titanian atmosphere is exposed to a significant CR flux that penetrates further into the atmosphere than on Ultima Thule. Cosmic rays thus create a second, lower ionosphere in Titan around 50–100 km altitude, which is comparable in intensity to the higher zone created by solar UV and trapped electrons. This CR-induced ionization of the Titan atmosphere has been modeled by several authors (Capone et al., 1980; Borucki et al., 1987, 2006; Molina-Cuberos et al., 1999a, 1999b), but of greatest significance in the interests of astrobiology is the role this ionization-driven atmospheric chemistry plays in creating organic molecules.

The upper, solar UV-generated ionosphere has been observed (by instruments aboard Cassini) to contain both positively charged organic molecules (Cravens et al., 2006) and heavy negative ions (Coates et al., 2007; Vuitton et al., 2009), which are believed to be important in the atmospheric hydrocarbon chemistry and transformation of simple gaseous species into organic-rich aerosol compounds—the so-called tholins. These long-chain hydrocarbons form the atmospheric haze layers and fall to deliver organics to the surface of Titan (Waite et al., 2007). Although these orbital detections of large species are from the upper ionosphere, similar organics are also expected to be produced in the lower, CR-generated ionosphere, and so contribute to the organic inventory and biological potential of Titan’s surface environment. Cosmic rays that penetrate through the atmosphere may also induce additional organic synthesis on the surface, especially in the polar lakes of methane and ethane (Raulin et al., 2009), such as driving the polymerization of HCN (Raulin et al., 1995).

While the terrestrial biosphere is largely sheltered from CR by the geomagnetic field and shielding atmospheric column, the surface of two planetary bodies with astrobiological potential (Mars and Europa), and meteorites potentially bearing microbial life ejected from planets receive no such protection. Indeed, the unshielded flux of CR onto the martian surface is one of the major environmental hazards to remnant surface life, even if present as spores or cryopreserved in permafrost.

7.8. Sterilization of Mars’ surface

Extensive geomorphological (Squyres and Kasting, 1994) and ground-based mineralogical and geochemical (e.g., Squyres et al., 2004) evidence indicates the ancient existence of flowing and standing water on the martian surface: a warmer, wetter Mars. The current martian surface environment is characterized by low temperatures, low atmospheric pressure, and consequent low availability of liquid water, and it is exceedingly hostile to the persistence of life (Horneck, 2000). The thin atmospheric layer offers only minimal shielding against solar UV (Patel et al., 2002), which would rapidly kill any exposed microorganisms (Schuerger et al., 2006) and photolyze organic molecules such as amino acids (ten Kate et al., 2005). This long-term UV flux is also expected to have produced a wind-dispersed layer of chemical oxidants on the surface (Zent and McKay, 1994; Yen et al., 2000).

Another major consequence of the thin atmospheric column and the absence of a martian dipole magnetic field (although it is believed to have once generated a global field; Acuña et al., 1999) is the poor shielding of Mars against CR. Energetic GCR and SEP are essentially unimpeded onto the martian surface, and the cascades of secondary particles they create penetrate several meters underground. Thus, the penetration of cosmic radiation probably exceeds the depth of oxidizing conditions in most parts of the martian surface and is a major threat to the survival of life held dormant by the freezing dry conditions.

A major research effort has been focused on modeling the martian surface ionizing radiation environment for the implications to human health during future astronaut exploration (e.g., Simonsen et al., 1990, 1991; Simonsen and Nealy, 1993; Kim et al., 1998; Cucinotta et al., 2002; De Angelis et al., 2004; Saganti et al., 2004). More specifically in the interests of astrophysics, several modeling studies have involved calculation of the radiation dose profile generated by CR through the martian regolith to predict the likely survival times of spore-forming or radiation-resistant bacterial strains such as D. radiodurans in the subsurface (Mileikowsky et al., 2000; Pavlov et al., 2002; Kninek et al., 2003; Dartnell et al., 2007a, 2007b, 2010). While there are uncertainties in both the physics of high-energy particle transport and the biological response to irradiation, these studies broadly agree and have found that the top 20 cm of the martian surface is sterilized of even the most radiation-resistant microorganisms (based on terrestrial models such as D. radiodurans) within a million years or so, with greater survival found at increasing depths due to shielding, especially in ice deposits (Dartnell et al., 2007a). In material with a density of around 1 g/cm³, the more energetic GCR cascades begin to dominate over the average SEP flux within around 10 cm depth (Dartnell et al., 2007a), and radiation from the GCR cascades peaks at (the Pfotzer maximum) 25–50 cm deep and at a third of that depth within solid rock or regolith (Mileikowsky et al., 2000; Pavlov et al., 2002; Dartnell et al., 2007b). Beneath the penetration of GCR cascades, at 3–4 m depth, depending on surface properties, the remaining source of ionizing radiation is the decay of radionuclides in the surrounding regolith. This natural radioactivity of martian regolith is believed to be much lower than that of terrestrial rocks (Mileikowsky et al., 2000) but would still sterilize deep dormant bacteria in 40 million years (Pavlov et al., 2002; Dartnell et al., 2007a).
While the flux of cosmic ionizing radiation onto the surface of Mars is likely to pose a considerable hazard to the survival of microbial life in the near subsurface, the situation may be reversed for Jupiter’s icy moon Europa, and ionization from particle radiation may actually help support a subsurface biosphere.

7.9. Support of europan biosphere

Europa is thought to possess a global ocean of liquid water beneath an outer ice shell of perhaps only ~1 km thickness (Greenberg et al., 2000), although other interpretations of surface features imply a much thicker (tens of kilometers) ice layer (Head et al., 1999). Sunlight cannot penetrate kilometers of ice, so a europan biosphere would not have access to widespread photosynthesis (Reynolds et al., 1983; Chyba and Hand, 2001). There has been some discussion, however, as to whether geothermal blackbody light emitted by seafloor hot water plumes (emission peaking at IR wavelengths) could support photosynthetic life, alongside chemoautotrophic organisms that harness inorganic chemical energy, around black smokers on Earth and, by extension, possibly Europa (e.g., White et al., 2002; Beatty et al., 2005; Rathgeber et al., 2008).

An alternative energy source for a chemosynthetic-based ecosystem could be provided by the charged particle radiation trapped within Jupiter’s magnetosphere (see Fig. 1), the jovian equivalent of Earth’s Van Allen belts (Chyba and Phillips, 2001). Particle radiation delivers more energy relevant to ice chemistry to the europan surface than solar UV and is comparable to internal heat flux from radiogenic or tidal heating (Cooper et al., 2001).

Irradiation of the outer surface water ice would produce oxidants, such as oxygen and hydrogen peroxide, and radiolysis of carbon dioxide in the ice and subsequent free-radical chemistry is expected to also generate reduced organic molecules such as formaldehyde (Delitsky and Lane, 1998; Chyba and Phillips, 2001). Such a radiolytically generated disequilibrium of oxidized and reduced chemical species could power a chemoautotrophic-based biosphere if these redox couples can be washed into the europan ocean, as discussed further below.

Niches of chemoautotrophs have been discovered on Earth that survive on radiogenic products. Lin et al. (2006) described a microbial community that survived at 2.8 km depth in basaltic crust, independent of photosynthesis or its products for tens of millions of years, by metabolizing sulfate with hydrogen released by radiolysis of the fracture water, driven by radionuclide decay in the surrounding basalt. There has even been a report of melanin-containing fungi, isolated from high-radiation environments like Chernobyl or the cooling water of nuclear reactors, that exhibit faster growth under conditions of high ionizing radiation flux than without it (Dadachova et al., 2007). The suggestion is that melanin enables the cells to extract energy from the ionizing radiation to drive their metabolism, a process termed “radiosynthesis.” A similar, weakly supported claim has been made for growth of photosynthetic bacteria and algae in the dark, supported by a low flux of gamma radiation (Luckey, 2008). If these preliminary studies on metabolism directly powered by ionizing radiation were to gather further corroboration, they would represent very interesting results indeed.

Thus, terrestrial chemoautotrophic organisms that are powered by nuclear decay energy are known, and a europan biosphere could be similarly maintained by redox couples generated by the ionizing flux of trapped particle radiation onto the surface ice. A radiogenically supported europan biosphere is only feasible, however, if these redox couples can be transported below to the ocean at high enough rate.

In an initial modeling effort, Chyba (2000a, 2000b) calculated that the rate of radiogenic nutrients imported into the ocean could support only a severely limited biomass. An updated model that incorporates new estimates of the depths of radiation processing and impact gardening of the europan surface ice has given a more optimistic conclusion (Chyba and Phillips, 2001). However, even assuming 100% conversion efficiency of this redox energy into biomass, microbial cell densities would still be very low and may frustrate attempts at biosignature detection.

While surface irradiation could maintain a microbial ecosystem, Chyba and Phillips (2001) deemed it unlikely that oxygen levels could rise high enough to support macrofauna akin to the aquatic animal life of our own oceans. On the other hand, Irwin and Schulze-Makuch (2003) drew attention to terrestrial invertebrates and vertebrates that can survive cold anoxic conditions and argued that metazoan life at the top of a europan food chain could perhaps approach the size of a tadpole or brine shrimp. More recently, Greenberg (2010) argued that various resurfacing mechanisms on Europa ensure the radiation-oxidized ice layer is at least 300 m thick, so the ocean could be oxygenated enough to support even a large macrofaunal ecosystem, of the order of 3 million tons of fish.

Another astrobiological consideration of CR is that the cellular damage they wreak may also pose an upper limit on the survival time of microorganisms embedded within meteorites and so restrict panspermia and the natural transfer of life between planetary bodies.

7.10. Panspermia

The theory of (litho)panspermia, that life can be transferred between planetary bodies within meteorites, has matured over recent years (see up-to-date reviews by Burchell, 2004, and Nicholson, 2009) with the results of simulations of planetary ejection by low-angle impact (Nyquist, 1983) or spallation (Melosh, 1984), orbital transfer dynamics calculations (Gladman et al., 2005), and experimental work on microbial survival of the associated shock pressures and temperatures (Fajardo-Cavazos et al., 2009), and resistance to the space environment (Horneck et al., 1994, 2001). On the whole, such experiments find that non-negligible proportions of microbial populations could indeed survive the shock of ejection or reentry and if protected from solar UV can tolerate the vacuum and desiccation of the space environment for protracted periods, thus supporting the conjecture of biological transfer within the Solar System. The possibility has also been discussed that Earth reseeded itself with returning ejecta after sterilizing impacts during the Late Heavy Bombardment (Wells et al., 2003).

During the transit between planetary bodies, microorganisms contained in the ejected rock fragments would be unprotected by planetary magnetic field or atmosphere, so
they would be exposed to the SEP and GCR flux, shielded from the charged particle bombardment only by the surrounding meteorite. Several studies have modeled the propagation of energetic CR particles through meteoritic rock to estimate microbial survival times. Mileikowsky et al. (2000) found that with meteorites smaller than about 60 cm in diameter the dose rate experienced in the center is actually increased by the shielding rock, due to the generation of secondary radiation cascades. For microorganisms sheltered in the center of a boulder at least 2.6 m in diameter, the shielding against external CR would be effective, and the radiation hazard from decay of radionuclides in the surrounding rock would begin to dominate. Using similar methods, Mileikowsky et al. (2000) and Clark et al. (1999) calculated that at 1 m depth in ejected regolith a sterilizing dose is accumulated on the order of millions of years. Mileikowsky et al. (2000) concluded that, within the distribution of likely meteorite sizes and transit times, the transfer of viable cells from Mars to Earth must be considered not only possible but highly probable. In the case of the martian meteorites so far recovered, however, Clark (2001) calculated that, due to their small size and period of interplanetary transit, all would have been exposed to strongly sterilizing doses of CR.

Interstellar panspermia, on the other hand, does not seem likely, considering the low probabilities of ejection from one planetary system and capture by another, the unlikelihood of safe landing on a planetary body at interstellar transfer velocity (Melosh, 2004), and the high accumulated levels of CR damage during the long transit. Valtonen et al. (2009) concluded that no potentially life-bearing ejecta from another solar system landed on Earth before life had already arisen on the planet, unless this transfer occurred between sister planetary systems in the same star cluster shortly after formation. Indeed, the accumulated background irradiation from GCR has been proposed as a natural mechanism for sterilizing (in the interests of planetary protection) any interstellar probes we launch, particularly to investigate exoplanetary systems (Cockell, 2008).

7.11. Degradation of biosignatures

As well as posing a survival threat to microbes in the martian surface or within meteorites, ionization and nuclear reactions driven by CR could degrade biosignatures indicating the prior existence of life (that may now have fallen extinct). Also, atmospheric ionization in a planet orbiting in the habitable zone of a M-class red dwarf star may act to mask the atmospheric disequilibrium indicative of a biosphere.

Considering Mars first, both experimental irradiation and radiation modeling studies have looked at the likely rate of erasure of different detectable biosignatures in the near subsurface. These include the destruction rate of amino acids (Kminek and Bada, 2006), racemization of an amino acid enantiomeric excess (Bonner, 2000, and references therein), alteration of biogenic organics (Court et al., 2006), nuclear reactions that lead to the loss of $^{13}$C/$^{12}$C isotopic bias (Pavlov et al., 2002), and biopolymer fragmentation (Pavlov et al., 2002). The general issue is that after millennia of cosmic irradiation it may be difficult to distinguish between breakdown debris of unambiguous biomolecules, a valid signature of extinct life, and simple abiotic organic species created in situ by prebiotic chemistry or exogenously delivered by meteoritic or cometary infall. There may only be a limited window of opportunity to find evidence of extinct surface life before it is degraded below the detection limit. Pavlov et al. (2002) estimated that in the martian near subsurface complete radiolytic macromolecular breakdown occurs within $10^7$ to $10^9$ years and eradication of a biogenic $^{13}$C/$^{12}$C isotopic bias in less than 25 million years. Kminek and Bada (2006) calculated that, to detect the remnant amino acids from Hesperian life, which occurred about 3 billion years ago, it would be necessary to drill to a depth of 1.5–2 m. However, the dependence of biomolecule destruction on radiation dose, particle linear energy transfer (see, e.g., Nelson, 2003), and temperature is not well understood (see discussion by Dartnell et al., 2007b), so there is uncertainty in these estimates.

Furthermore, the biosignature detection instrumentation itself may be vulnerable to ionization by CR, particularly proposed designs that employ organic dyes or antibodies, and this possibility has been explored recently. Thompson et al. (2006) ran proton and alpha irradiations on fluorescent dyes proposed for use in labeled assays, and Le Postollec et al. (2009a, 2009b) conducted both modeling and experimental neutron irradiation work on antibodies and dyes. These studies found that the antibody function or fluorescence response was not significantly affected by ionizing radiation fluences comparable to those expected during a martian mission.

Cosmic rays may also play a significant role in masking biosignatures in the atmospheres of extrasolar planets. Biosignatures could potentially be remotely detected on exoplanets by IR spectroscopy and identification of atmospheric chemistry far from thermodynamic equilibrium, such as the co-presence of oxygen (or its photolytic product, ozone) and methane (Lovelock, 1965; Sagan et al., 1993; Des Marais et al., 2002; Kaltenegger et al., 2007). Planets orbiting within the habitable zone of M-class stars are expected to possess weaker magnetospheric shields and so would be exposed to higher fluxes of GCR than Earth, as discussed above in Section 7.2. Grenfell et al. (2007) considered the likely effects of this more intense GCR environment in producing nitrogen oxides and driving atmospheric chemistry to modify potential biomarker molecules such as ozone, nitrous oxide, and methane, all of which can reach much higher concentrations in the atmosphere of M-class star planets and make particularly conspicuous biosignatures (Segura et al., 2005). They found that the profile of ozone in the planetary atmosphere is only modestly affected by GCR-induced nitrogen oxide production, while nitrous oxide and methane show negligible difference. Thus, this is not anticipated to hinder spectroscopic detection of biosignatures on planets in the habitable zone of M-class stars. However, as with Griebmeier et al. (2005), Grenfell et al. (2007) did not consider the SEP flux from the host star itself, which may constitute a non-negligible atmospheric ionization source. Segura et al. (2010), whose work will be discussed further in Section 8.3, did model the effects of SEP flux from a particularly large flare on an active M dwarf hitting an Earth-like planet that lacked a magnetic field in the habitable zone. They concluded that, although atmospheric ozone may be severely depleted, methane concentration is not perturbed, and so a single large
flare would not hinder spectroscopic atmospheric bio-
signature detection. The effects of repeated flares before the
atmosphere can return to equilibrium, however, are not well
understood.

At the beginning of this section, I reviewed how the
background flux of CR onto Earth, or any similar planet
shielded by magnetic field and thick atmosphere, poses no
significant threat to life on the surface. However, extreme
astrophysical events can release an intense flux of CR or
energetic electromagnetic radiation and, thus, present a
sporadic but acute ionizing radiation hazard to a planetary
biosphere, as is discussed in detail in the next section.

8. Extreme Radiation Events

While a planetary magnetic field or significant atmo-
spheric column can effectively shield the surface of a habit-
able world from the radiological effects of the background
CR flux, even an Earth-like planet could be seriously affected
by a more intense pulse of ionizing radiation. The conse-
quences of an extremely energetic particle event, with a high
enough fluence and average particle energy to penetrate
even dense atmospheric shields and thus present a direct
radiation hazard on the surface, should not be neglected.

Dar et al. (1998) discussed the effects of a high-energy CR
jet that may also be generated at the same time as a GRB
(Vietri et al., 2003). Dar et al. (1998) argued that such a highly
collimated relativistic particle jet would have a beaming
angle $\Delta \Omega \leq 0.01$, similar to those observed or estimated for
active galactic nuclei and microquasars, and would reach a
distance of around 1 kpc before becoming disrupted. If such
a jet were to strike a terrestrial planet, the particle flux would
not only be extremely high but the average particle energy
would be a thousand times higher than the background CR.

The energy deposition in the atmosphere over days to
months would equal that of the normal background GCR
over 10 million years. A large fraction of the energy of the jet
primary particles, with a typical energy of $\sim 1$ TeV/nucleon,
would be transformed into energetic muons by extensive air
showers. These are highly penetrating and would irradiate
with lethal doses to depths of hundreds of meters of water or
rock. In addition to the ionizing muon flux and stratospheric
ozone depletion, the particle jet would create radioactive
isotopes in the atmosphere and surface by nuclear spallation,
which would be dispersed globally by winds. Dar et al.
(1998) estimated that such a CR jet might be targeted toward
Earth by a nearby neutron star merger on the order of every
100 million years and so could account for some of the five
metazoan mass extinctions of the past 570 million years.

Even if energetic CR do not penetrate to a planetary sur-
face to pose a direct radiological hazard, there are significant
consequences of the extreme atmospheric ionization caused
by violent astrophysical events like supernovae or GRB. The
potential outcomes of extreme ionization events are sum-
mORIZED in the schematic diagram in Fig. 7, with the labels
relating to the main text section numbers below.

8.1. Supernova destruction of the ozone shield

The danger presented to a biosphere by intense photon or
particle radiation from a source such as a supernova, even if
no significant amount of the ionizing radiation directly
reaches the surface, was recognized by Ruderman (1974).
Ionizing radiation propagating through the atmosphere
causes the dissociation of N$_2$ to produce NO$_x$ compounds
(most crucial NO and NO$_2$), which act catalytically to de-
stroy stratospheric O$_3$. The ozone layer shielding Earth’s
surface from biologically active bands of solar UV radiation
is only 3–4 mm thick at standard temperature and pressure,
and Ruderman (1974) estimated that this shield would be
substantially depleted for up to a century by a supernova at
50 light-year range. UV is harmful to both DNA and pho-
synthetic pigments (Horneck, 1995), and an increase in flux
of only 10–20% can be lethal to many terrestrial organisms
such as phytoplankton, the base of many food chains
(Thomas and Melott, 2006). Aikin et al. (1980) argued that
stratospheric ozone would actually be enhanced at first by
the intense UV flash of a nearby supernova, before the
gamma-ray flux overtakes Earth several months later and
destroys the ozone again.

In subsequent years, many authors have attempted to
quantify the degree of stratospheric ozone depletion and
time before recovery, and likely effects to the biosphere, from
a nearby supernova explosion (e.g., Ruderman, 1974; Whitten
et al., 1976; Reid et al., 1978; Ellis and Schramm, 1995;
Crutzen and Brühl, 1996; Gehrels et al., 2003; Thomas and
Honeyman, 2008). The modeled results often disagree;
therefore, the conclusions drawn disagree as well with re-
gard to the extent of biological perturbations on terrestrial
planets from nearby supernovae. Recently, for example,
Gehrels et al. (2003) considered ozone depletion from both
the gamma-ray photons and later-arriving pulse of CR from
a supernova in a two-dimensional atmospheric model. They
calculated that, for the surface biologically active UV flux to
be doubled, the supernova must occur within 8 pc of Earth,
an event which happens only every 1.5 billion years; thus
Gehrels et al. (2003) concluded that supernovae are less im-
portant than other potential causes of mass extinction. In
general, impacts on Earth’s biosphere from nearby super-
novae are estimated to occur on the timescale of every few
hundred million to several billion years; thus they do pose a
plausible hazard to life. There is no evidence, however, for a
direct link between supernovae and a past mass extinction.

Thomas et al. (2008) considered the likely effects on Earth of
a specific massive star, $\eta$ Carinae, that appears ready to
detonate. At 2.3 kpc distance, $\eta$ Carinae is unlikely to pose
any serious hazard as a normal supernova, but it may pro-
duce a superluminous supernova that could have potential
implications even at such a great range. However, Thomas
et al. considered atmospheric ionization (and subsequent
ozone depletion), enhanced CR flux, and endocrine disrup-
tion from intense blue light, and concluded that even if su-
perluminous, $\eta$ Carinae would be unlikely to endanger the
terrestrial biosphere.

8.2. Gamma-ray bursts

Gamma-ray bursts are far more energetic than supernovae
and have been identified as a potential threat to biospheres
by Thorsett (1995). Thorsett (1995) estimated that GRB,
which are potentially caused by the merging of binary neu-
tron stars, occur in our own Galaxy about every $10^7$ to $10^8$
years, with a burst occurring uncomfortably close (within
1 kpc) perhaps every $10^5$ years (see also Melott et al., 2004,
and Dermer and Holmes, 2005, for more recent estimates of the frequency of GRB per galaxy). Thorsett (1995) calculated that a GRB as distant as the galactic center, almost 8 kpc remote, would have similar consequences to the nearby supernova considered by Ruderman (1974). It has also been suggested that the occurrence of GRB could periodically sterilize the Galaxy of land-based life (Annis, 1999).

The effects of supernovae or GRB on a planetary atmosphere or biosphere would be similar, although from very different effective ranges (see Fig. 7). Considering both their relative intrinsic intensity and galactic frequency, the greatest danger to a biosphere is posed by supernovae less than 10 pc distant (Reid et al., 1978; Ellis and Schramm, 1995; Gehrels et al., 2003) and directed GRB 1–100 kpc remote (Scalo and Wheeler, 2002; Melott et al., 2004, 2005; Dermer and Holmes, 2005; Thomas and Melott, 2006; Martín et al., 2009).

Aside from intrinsic luminosity, the major difference between GRB and supernova events is the time course over which the pulse of energetic photons occurs. Ejzak et al. (2007) studied how the atmospheric end point of extreme ionization events from GRB or supernovae is affected by the blast duration of ionizing photons (from $10^{-1}$ to $10^8$ s). They found that the ultimate atmospheric effects do not vary significantly, so the effects of extreme atmospheric ionization from different astrophysical events can be satisfactorily approximated by disregarding the time development of such events and considering only the total radiation fluence.

Analyzing the effects of a GRB within 2 kpc of Earth, Thomas et al. (2005a) calculated that the irradiation would globally deplete stratospheric ozone by 35%, an event that is statistically estimated to have occurred at least once in the last billion years. Ozone recovery from such destruction would take at least 5 years, during which time the surface would be exposed to up to a 3-fold increase in solar UVB flux, which the authors predict would trigger widespread extinction.

The threat posed to terrestrial planetary atmospheres and biospheres by GRB has been reviewed recently by Thomas and Melott (2006). They noted that, although polar regions experience greatest long-term ozone depletion and

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**FIG. 7.** Schematic summarizing the cascade of effects and hazards to the biosphere of extreme ionizing radiation events imposed on a habitable planet by sporadic astrophysical sources. The labels refer to subsection numbers in Section 8 of the main text. Color images available online at www.liebertonline.com/ast
consequent enhanced UVB flux, DNA damage is limited here by reduced incident solar irradiance, so the increase in UVB flux is most hazardous to land-based or shallow water-borne life (less than 10 m in depth) in midlatitudes.

Galante and Horvath (2007a, 2007b) split the impact of a GRB-directed beam into four distinct phenomena: direct gamma-ray transmission to the planetary surface, the UV flash created by reemission in the atmosphere (an effect first considered by Smith et al., 2004a; see Section 8.6 of this current review), ozone depletion, and increased flux of CR. They found that the first three can affect a biosphere from huge distances, greater than the diameter of the Galaxy, whereas the enhanced CR flux would only be lethal for a very close GRB, only a few parsecs away.

8.3. Superflares

Extreme ionization of the atmosphere could also be caused by a source much more local than GRB or supernovae: an anomalously energetic solar particle event. Reid et al. (1976) considered the likely terrestrial effects of solar flares 10 and 100 times more intense than the August 1972 event, widely accepted to be the largest solar particle event of the space age (Mewaldt, 2006). They calculated a global ozone depletion of around 25% and 50%, respectively, and a consequent increase in biologically effective UV flux at the surface by a factor of 1.5 and 2.5, but only if such a flare coincidentally occurred within the period of a few thousand years during a polarity reversal of the geomagnetic field. Even without the co-occurrence of an intense solar particle event, Reid et al. (1976) calculated stratospheric ozone levels would drop by 4.3% during polarity reversal due to increased penetration of GCR.

Thomas et al. (2007) modeled the specific case of the 1859 “Carrington event” solar flare, which is believed to be one of the most powerful white-light flares ever observed. Nitrate deposits in polar ice indicate that the Carrington flare was the largest solar proton event of the last 450 years, with a >30 MeV solar proton fluence around 4 times greater than the measured August 1972 event (Shea et al., 2006). Thomas et al. (2007) calculated that the intense proton flux from such a flare would have caused up to 14% localized ozone depletion for about 4 years and an increase in nitrate deposition such that it would agree reasonably with measurements taken from ice cores. On the other hand, Jackman et al. (2000) modeled the perturbation to terrestrial stratospheric chemistry and ozone destruction due to two large solar particle events from the past 30 years and modeled the August 1972 and October 1989 events as well. They estimated that the chemical state of the atmosphere, such as stratospheric chlorine levels, can actually cause an increase in total ozone for a few years after such an event.

The hazard from flares may be particularly acute for Earth-like planets orbiting within the habitable zone of active M-class red dwarf stars. Segura et al. (2010) modeled the atmospheric effects of the UV and SEP flux from a large flare hitting an Earth-like planet that lacked a magnetic field in the habitable zone of the active M dwarf AD Leonis. They found that the UV emission from the flare had negligible impact on atmospheric ozone but that the nitrogen oxides produced by ionization from the energetic protons resulted in an ozone depletion of more than 90% within two years of the flare, which took almost 50 years to recover. However, even at maximum ozone depletion, the surface UV dose rate for DNA damage was only 4% larger than that on Earth.

Alongside the destruction of the ozone shield, several other potentially catastrophic insults to a planetary biosphere from an acute atmospheric ionization event have been identified, which are reviewed in Subsections 8.4–8.6 below.

8.4. Global cooling

Beyond ozone destruction, the elevated levels of nitrogen dioxide generated in the atmosphere by cosmic radiation from a superflare, supernova, or GRB are also expected to cause other effects on planetary habitability. Nitrogen dioxide is a brown gas that absorbs strongly at visible wavelengths, so a spike in atmospheric NO$_2$ levels could significantly reduce solar insolation that reaches the ground and could potentially trigger global cooling.

Stratospheric nitrogen dioxide produced by the superflare modeled by Reid et al. (1976) was calculated to decrease surface visible solar irradiance by several percent, but due to the heat reservoir of the oceans such a transient perturbation to the global thermal budget may not trigger any lasting climatic change (Reid et al., 1978). For a nearby supernova, Reid et al. (1978) postulated that the resultant cooling of Earth’s surface (estimated at up to a 3 K decrease in global average temperature) from nitrogen dioxide opacity may in fact cause a more pronounced biological impact than the ozone destruction and increased surface UV flux. Such a global cooling event would reduce atmospheric water vapor content by about 20%, which would result in further cooling through reduced efficacy of the greenhouse effect and widespread drought conditions. Melott et al. (2005) found that, averaged globally, atmospheric NO$_2$ produced by a GRB would reduce surface solar irradiance by 1%, but up to 30% reduction in polar regions (due to poleward transport of GRB products within the stratosphere), for several years. Melott et al. (2005) speculated that such an input could possibly trigger a global cooling event, although Thomas reported that no detailed simulations of the likely climatic response have yet been run (Thomas, 2009).

Tanaka (2006) considered another potential mechanism for triggering global cooling by a nearby supernova: greater CR flux and ionization of the lower atmosphere. These atmospheric ions would act as condensation nuclei and thus increase cloud formation at low altitude, possibly triggering a major ice-age epoch.

A further complication of the likely climatic consequences is that nitrogen dioxide shading and cooling of the lower atmosphere and surface is not equivalent to a simple decrease in solar irradiance, as this absorbed energy leads to heating of the stratosphere. A hotter stratosphere with increased thermal emission to the ground would partially compensate for visible shading, which would reduce surface cooling by up to a factor of 2 (Reid et al., 1978; Melott et al., 2005) but also affect the dynamic interaction between the stratosphere and troposphere with possible consequences for global climate.

8.5. Nitric acid rain-out

As the atmosphere returns to its pre-perturbation composition, nitrogen oxides are removed by oxidation and
Thus, even a short-duration burst of energetic cosmic radiation can have effects on atmospheric chemistry and, consequently, the biosphere, that are long-lived and global in their impact.

8.6. Ultraviolet flash

In addition to the above delayed threats from extreme atmospheric ionization, shielding of an intense CR pulse by a planetary atmosphere can have other deleterious effects to a surface biosphere, only recently described. Over the past few years, Smith, Scalo, and Wheeler have co-authored a series of papers detailing their investigations into the transformation of the incident cosmic ionizing radiation by the shielding atmosphere into biologically hazardous bands of UV light, through a process similar to aurora: electronic excitation of atmospheric molecules and reemission of photons.

Smith et al. (2004a) considered the flux of cosmic X-rays and gamma rays from sources such as solar flares, supernovae, and GRB, incident onto habitable exoplanets with varying atmospheres. Even thin atmospheres (although still dense enough to maintain habitable conditions on the surface) will block incident X-rays, but still transmit a significant fraction of the gamma-ray flux to the surface. Exoplanets with atmospheres thicker than about 100 g/cm$^2$ would be effectively shielded against the penetration of both X-rays and gamma rays, but interactions within the atmosphere would reprocess almost all the incident photon energy into diffuse UV emission. These longer wavelengths are more effectively transmitted through the atmosphere (although dependent on atmospheric thickness and UV absorbers), and for a planet with an atmosphere of around 100 g/cm$^2$ up to 10% of the incident energy could reach the planetary surface as UV reemission from the atmosphere, which would subject exposed surface organisms to a harsh UV insult.

Smith et al. (2004b) expanded on this analysis for Archean Earth (3.9–2.5 Ga), considering the photon spectrum from nearby supernovae as well as solar flares. Assuming only Rayleigh scattering of UV photons by the $\sim 1000$ g/cm$^2$ atmosphere and no UV filters such as O$_3$, Smith et al. (2004b) found that up to 4% of the incident ionizing radiation reaches the surface as biologically effective UV radiation (200–320 nm; UVC and UVB bands). For the current terrestrial situation, including an ozone shield, only 0.2% of the incident energy reaches the surface as biologically active UV reemission, although this is still many orders of magnitude higher than the energy of the ionizing radiation that penetrates to the surface. Considering just the solar flare influence on planets orbiting within the habitable zone of a Sun-type star, significant fluxes of hazardous UV would occasionally penetrate to the surface, even on evolved terrestrial planets with a photosynthetic biosphere and ozone shield, but would not exceed the direct UV from the sun. For terrestrial planets orbiting within the habitable zone of lower-mass stars (<0.5 solar mass), such as M-class dwarves, however, the energy from solar flare X-rays and gamma rays redistributed into short wavelength UV by the atmospheric shield would completely dominate the direct UV flux from the parent star (Smith et al., 2004b). Organisms evolving on such a planetary surface may therefore experience frequent pulses of elevated UV flux and, if not adequately protected, consequent episodes of cellular damage or hypermutation.
Scalo and Wheeler (2002) attempted to calculate the frequency with which terrestrial planets are exposed to an ionizing radiation flux from a GRB in the Galaxy sufficient to cause a direct biological effect through DNA damage. About 100–500 such potentially biologically significant events or “jolts” are predicted to occur every billion years for eukaryotic-like organisms and around 100 times fewer for prokaryotic life, which is more radiation resistant, through direct ionizing radiation beneath thin atmospheres (such as modern-day Mars) or through UV reemission by thick atmospheres (such as Earth). However, due to the short duration of GRB (only one hemisphere of a planet would be exposed) and the effective shielding of organisms provided by only 50 cm of water or shallower depths for rock, only a tiny proportion of these stochastic irradiation events might actually significantly perturb the biosphere.

Martin et al. (2009) also considered the UV flash generated by atmospheric reemission from a GRB at different epochs in the evolution of Earth’s biosphere and atmosphere, and they reported that a surface ecosystem during the Mid-Proterozoic would likely be most stressed. Galante and Horvath (2007a, b) presented a more comprehensive analysis of GRB impact on an Earth-like atmosphere and biosphere, modeling the effect of direct gamma-ray transmission, UV flash, stratospheric ozone destruction, and enhanced CR flux. They found that global environmental and biological upsets may be caused by GRB sources up to 100 kpc remote. Peñate et al. (2010) found that the UV reemission from a GRB 2 kpc away could penetrate up to 75 m through clear water (20–30 m for more turbid water) to severely affect aquatic photosynthetic primary producers, particularly if a night flash, to a great depth, and thus also affect the ecosystems depending upon them.

In any case, it has become clear that the pulse of ionizing radiation from a supernova or GRB can affect the planetary surface UV environment (beneath an atmosphere) in both the short and long term. Atmospheric irradiation and reemission produce a transient burst of UV, while the perturbed atmospheric chemistry and nitric oxide generation leads to destruction of an ozone shield and increased transmission of solar UV. The rest of this section deals with variations in the GCR flux incident on a habitable planet.

### 8.7. Active galactic nuclei

Another astrophysical source of energetic ionizing radiation, alongside supernovae, GRB, and superflares, which is not so often discussed, is the accelerated particle emission from the active nucleus of a galaxy. The energy source driving active galactic nuclei is believed to be a supermassive black hole and the surrounding accretion disc; such objects are very luminous in X-rays. If the core of the Milky Way has been active in this way in its past, or in other galaxies during a phase of nucleus activity, there would also be an enhanced particle radiation flux that is likely to be a more significant hazard to habitable systems in the galaxy than the X-ray flux. Clarke (1981) considered a simple model of the situation and found that the particle radiation flux at Earth would be enhanced by a factor of around 100 during the lifetime of galactic nucleus activity. Such an increased ionizing particle flux would be concomitant with planetary hazards such as ozone depletion and enhanced surface radiation environment as described above.

### 8.8. Variation in the galactic cosmic ray flux

Beyond possible activity of the core of the Milky Way, other mechanisms may cause variation in the GCR flux incident on the Solar System. This variability of the incident GCR flux is independent of the modulation of the GCR flux through the Solar System by the varying heliosphere (see Section 7.1, as well as reviews by Scherer et al., 2006, and Vainio et al., 2009). Although diffusion of charged particles along interstellar magnetic field lines acts to produce an essentially isotropic and uniform background flux of CR throughout the Galaxy, moderate statistical variations are expected due to the random space-time distribution of supernovae (Erlkyin and Wolfendale, 2001). In particular, Svensmark (2006) modeled the CR flux into Earth’s atmosphere over the history of the planet based on estimations of the galactic star formation (and thus also supernova) rate, SEP flux, and diminishing solar activity modulation of GCR, and reported that particle fluxes were at most twice their current intensity over the past 4.5 billion years. The distribution of supernovae throughout the Galaxy and, thus, the expected frequency of acute cosmic ionizing radiation events impacting a habitable planet is also a critical factor in defining the extent of the galactic habitable zone and its spatial evolution over time (Lineweaver et al., 2004).

Medvedev and Melott (2007) argued for a causal link between a claimed $62 \pm 3$ million-year cycle in Phanerzoic animal fossil diversity (discovered by Rohde and Muller, 2005, and confirmed to be a robust signal by Lieberman and Melott, 2007) and periodic variation in the GCR flux. They argued that, due to the vertical oscillation of the Solar System through the plane of the Galaxy (with a period of 64 million years) and an anisotropy in the extra-GCR flux caused by the motion of the Milky Way through the intergalactic medium, there is a cyclical variation in the CR impinging on Earth over the right timescale to explain the fossil diversity variation. Although Medvedev and Melott (2007) proffered no mechanism for actually linking extra-GCR flux and terrestrial biodiversity, Melott et al. (2008) investigated one plausible mechanism: CR-induced destruction of ozone and increased surface exposure to solar UV. Even at the upper end of the range of plausible parameters explored, Melott et al. (2008) found only a 2% reduction in global ozone levels, no more than currently being caused by anthropogenic pushing of atmospheric chemistry (such as chlorofluorocarbon release) and far less than that predicted for a nearby GRB. So even though the duration of increased atmospheric ionization and consequent ozone depletion from enhanced extra-GCR is predicted to last around a million times longer than that from a GRB, it is unlikely to exert a great biological stress unless coincidentally occurring in combination with other threats to the biosphere.

In addition to variation in the background GCR flux, the atmospheres and biospheres of planetary bodies are exposed to extreme stochastic events such as the superflares, nearby supernovae, and remote GRB discussed above. Thus, the cosmic ionizing radiation environment of a potentially habitable planet or moon is highly dynamic, with temporal variability across a great range of timescales:

- minutes for the most intense stage of GRB,
- hours to days for solar particle events,
• days to weeks for the initial electromagnetic pulse from supernovae,
• thousands of years if the solar system passes through the expanding shock cavity of a nearby supernova remnant or the heliosphere is compressed by passage through dense interstellar clouds,
• tens of millions of years from oscillations of the Solar System about the galactic plane.

Variations in the GCR flux, or indeed stochastic pulses of extreme ionizing radiation flux from nearby supernovae or GRB, may be recorded by spallation products and cosmogenic isotopes embedded in the lunar regolith and then shielded and protected by a later lava flow (Crawford et al., 2010; Fagents et al., 2010). Measurement of the CR bombardment history of unprotected planetary surfaces such as meteorites and lunar rocks was reviewed by Reedy et al. (1983) and Eugster (2003).

So far in this review, I have considered the effects of exogenous cosmic radiation incident onto a planetary atmosphere or surface. Potentially habitable planets and moons are also likely to include a significant component of radiogenic isotopes and so generate their own intrinsic ionizing radiation environment, the astrobiological implications of which are explored in the following section.

9. Intrinsic Planetary Radiation Environment

The decay of radioisotopes within a planetary body serves as a heat source (along with remnant heat from formation) for volcanism and plate tectonics, both of which are thought to have been important in maintaining habitable conditions on Earth for long periods (e.g., see Lammer et al., 2009). In addition to this role in habitability, the intrinsic radiation environment of a planet or moon, generated by radioisotope decay as opposed to CR impinging onto an atmosphere and biosphere, is also a critical factor in habitability and the origin of life. On Earth, ambient radiation levels from radiouclide decay peaked about four billion years ago, at a level approximately 7 times higher than today, and have been steadily decreasing since (Karam and Leslie, 1999). However, the rise in free oxygen over geological time and its enhancing effect on DNA damage through increasing reactive oxygen species generated by radiolysis may have served to steady mutation rates over the history of life (Karam et al., 2001). The role of enhanced primordial radioactivity in driving prebiotic chemistry on Earth is reviewed below.

9.1. Radiation and prebiotic chemistry

The radiolytically driven production of biologically relevant molecules in primordial Earth’s oceans, from decay of potassium-40, was considered by Draganic (2005). While the annual energy available from dissolved potassium-40 was much less than that from other sources, it is argued to be significant, as it was distributed evenly throughout the primordial ocean, including the deep ocean where other energy sources such as solar UV would have been absent. It is estimated that 10^15 g of organic compounds would have been created over 100 million years by in situ radiation in the primordial oceans (Draganic, 2005).

Primordial carbonate minerals, co-deposited with radiouclide metal cations such as potassium-40, rubidium-87, and the uranium-235, uranium-238, and thorium-232 radioactive decay families, would have self-irradiated and may have generated substantial quantities of organic precursor species (Collins et al., 2000). Production of organics by carbonaceous irradiation has been achieved experimentally (Kolomnikov et al., 1982; Albarrán et al., 1987), and on early Earth these precursors could have entered the hydrosphere and contributed to the origin of life by direct dissolution or weathering out of the carbonate rocks. On the other hand, synthesis of organic molecules on clay mineral substrates may have been hampered by potassium-40 radiolytic destruction exceeding the rate of free-radical–induced synthesis or polymerization (Draganic, 2000). There may have also been significant synthesis of prebiotic molecules at kilometers depth in the Archean terrestrial crust from radionuclide irradiation of water held in the porous rock (Garzón and Garzón, 2001).

It has been proposed that the observed enantiomer bias in terrestrial biology, the exclusive use of L-amino acids and Rsugars, was fostered on primordial Earth by the role of radiouclide decay in prebiotic chemistry (see Fitz et al., 2007, for a review of theories proposed to explain the origin of biological homochirality). For example, the polarization of electrons released by beta decay of radionuclides has been suggested as a causative agent in selecting the enantiomer bias by preferential synthesis or degradation of either enantiomer in an initially racemic mixture. However, systematic reviews of the varied experimental approaches attempted to demonstrate such an effect have concluded that there is no substantiating evidence for such a causal connection between nuclear parity violation and selection of an enantiomer bias in biomolecules (reviewed in Bonner, 2000, and Fitz et al., 2007). A more recent alternative hypothesis posits that supernovae in the vicinity of the ISM from which the Solar System condensed were responsible for the generation of the observed bias toward L-amino acids in carbonaceous chondrite meteorites. Boyd et al. (2010) argued that amino acid enantiomers become aligned in opposite directions along intense magnetic field lines and consequently R-amino acids are preferentially destroyed by the pulse of neutrino radiation from a nearby supernova. The mechanism proposed by Boyd et al. (2010) is that atoms of 14N would become aligned along the intense magnetic field lines of the nascent neutron star in a manner dependent on the electronic structure and, thus, chirality of the parent molecule; hence they would also force the orientation of amino acids or their precursors that contain the 14N atom. The subsequent pulse of electron antineutrinos from the supernova would preferentially destroy one spin orientation of the 14N and thus the chiral molecule containing it, which would produce an enantiomeric bias in the surviving molecules in the ISM out to a significant range from the supernova. Later chemical amplification of this enantiomer bias, and incorporation of the astrochemistry-produced organics into carbonaceous chondrites and delivery to the surface of primordial Earth, could explain the observed L-amino acid preference in the terrestrial biosphere.

Parnell (2004a) considered a subsequent role for radioactivity of mineral sands in the origin of life, as an energy source that drives the polymerization and precipitation of complex organic molecules from precursors. Methane outgassed from an impact crater through hydrothermal systems
containing radioactive minerals could also have been an important site for organic synthesis on early Earth or perhaps Mars (Parnell et al., 2006). Indeed, carbonaceous coatings are often observed as having accreted around uranium- or thorium-rich mineral grains of Archean sedimentary rocks. Court et al. (2006) experimentally observed ionizing radiation from uranium- and thorium-rich grains polymerizing simple hydrocarbons such as methane into complex polycyclic aromatic hydrocarbon–rich organic matter. They argued that the prebiotic organic inventory synthesized from such a process, either in situ on early Earth or delivered from irradiated extraterrestrial material, made a significant input to the origin of life.

Such radiolytically driven organic synthesis may even feed a limited biosphere on an orphaned terrestrial planet gravitationally ejected from its solar system into interstellar space (Stevenson, 1999).

9.2. Natural fission reactors

An example of an extremely intense intrinsic source of particle radiation on a planetary surface is provided by the ancient natural fission reactors in Oklo, Gabon (see reviews by Zetterström, 2000, and Jensen and Ewing, 2001). Here, multiple small pockets of uranium ore contained within layers of sedimentary rocks achieved criticality around 2 billion years ago. Lens-shaped reactor cores, around 10 m in diameter and 0.5 m thick, sustained nuclear fission on and off for 10⁷ to 10⁸ years (Nagy et al., 1991), with the water table acting as an efficient moderator of neutrons and the fission reaction (Gauthier-Lafaye et al., 1996). The fission reaction created temperatures estimated between 160°C and 360°C during criticality (Nagy et al., 1991), which drove the circulation of hot fluids and hydrothermally altered the host rock for up to 50 m from the reactor core, transforming the sandstone strata into clays (Gauthier-Lafaye et al., 1996) and producing abundant solidified bitumen from organic molecules originating in the host rock (Parnell, 1996).

Close to the reactor core, the radiation field would have been a complex mix of fission fragments, energetic electrons, neutrons, and gamma rays (Draganic et al., 1983). It has been estimated by in-pile dosimetry that, operating at a kilowatt-level power, the Oklo natural nuclear reactors would have generated radiation dose rates of over 1 kGy/h within their cores. However, a more constraining biological hazard would have been the high temperatures created by the fission reaction. With the thermal gradient in the clays surrounding the core calculated to be 100°C/m (Gauthier-Lafaye et al., 1996), microbial life several meters away from the cores may have been able to survive the thermal hazard and been exposed to only the more penetrating neutron and gamma radiation from the fission reactions. Furthermore, 14% of the fission products have half-lives longer than one year and so would have created an elevated radiation environment for subterranean life during periods of reactor shutdown and also migrated outward from the core into the surrounding environment (Draganic et al., 1983).

The Oklo reactors were active 2 billion years ago, which was too late for an involvement in the origins of life. Indeed, the conditions that favor the formation of concentrated deposits of uranium necessary to achieve criticality are thought to have been created by a redox boundary produced by biology (Gauthier-Lafaye and Weber, 2003). Uranium is soluble under slightly oxic conditions and when reduced it precipitates, usually as uranite. The Oklo reactors were created when the water-soluble U⁴⁺ ion, which was oxidized in the groundwater from dissolved oxygen released by early cyanobacteria, migrated underground until it encountered a reduction front created by subsurface organics. The uranium precipitated out of solution along this redox boundary to produce a rich deposit.

Abiotic processes for creating sufficiently enriched deposits of fissile isotopes, which could have played a role in life’s origin, have also been discussed. Adam (2007) presented the hypothesis of a possible role in prebiotic chemistry for a Hadean or early Archean heavy mineral placer beach 4.3 billion years ago. A placer beach is formed by gravitational separation of mineral grains by density during the sedimentation process and can concentrate heavy minerals like uranium to 70–90 wt% at particular sand horizons up the beach. Such an enrichment of heavy mineral grains, including radioactive isotopes of thorium and uranium (which would have been stable in the non-oxidizing primordial atmosphere) as well as other actinide elements, would have favored much prebiotic chemistry. Actinides can act as catalysts for oligomerization; and the energy released by their radioactive decay, and potentially by a self-sustaining fission reaction far earlier than the Oklo-type deposit discussed above, could have driven radiolysis and synthesis of simple organic molecules from inorganic carbon dissolved in the seawater between mineral grains, as Adam also demonstrated experimentally.

If synthesis and concentration of organic molecules by radioactive grains is indeed a necessary process in the origins of life, then this is another way in which plate tectonics may be a prerequisite for the habitability of a terrestrial planet (in addition to maintaining the composition of the atmosphere by recycling volatiles; e.g., see recent review by Lammer et al., 2009). The concentration of radioactive monazite and uranite mineral grains in certain locations is dependent on the erosion of granitic continental crust, so plate tectonics may perform a vital role in driving prebiotic chemistry (Parnell, 2004b).

10. Conclusions

This review has surveyed the broad range of current research where ionizing radiation from exogenous CR, intense electromagnetic radiation from sporadic astrophysical sources, and intrinsic mineral radioactivity of a world exerts significant influence over processes pertinent to astrobiology. Ionizing radiation has traditionally been considered as mainly detrimental to biology, and ionization and radiolysis are indeed greatly destructive to the complex molecules and polymers of life and thus pose a severe hazard to the development and survival of organisms. So, the ionizing radiation flux may well constitute a limiting factor for the persistence of life on unshielded planetary surfaces such as Mars (or perhaps high in the venusian atmosphere), on worlds orbiting within the habitable zone of stars less massive than the Sun, and during the interplanetary transfer of microorganisms via panspermia. Earth’s surface is protected from CR by screening from the Sun’s heliosphere, the global magnetosphere, and absorption of particle cascades by the
atmosphere, and any habitable exoplanet would experience similar multiple layers of shielding. However, even well-protected Earth-like worlds may receive biosphere-threatening radiation hazards during rare, but extreme, astrophysical events. Cosmic rays may also act to degrade or mask detectable biosignatures in the martian surface or atmosphere of an exoplanet, as well as impairing the functionality of the biodetection instrumentation on probes.

The synthesis presented here, though, has demonstrated the many varied ways in which ionizing radiation is also beneficial and plays a vital role in providing habitable environments and supporting the development of life in the Cosmos. These functions include the synthesis of biogenic elements and the driving of prebiotic chemistry to generate organic molecules in the ISM, circumstellar dust during formation of planetary systems, and the atmospheres, oceans, and subsurfaces of habitable planets and moons. Indeed, in the case of Europa, the flux of trapped ionizing radiation onto the icy surface and radiolytic production of oxidants and organics may even nourish an oceanic ecosystem; a curious counterpoint to the biosphere-limiting effect of CR in the martian near subsurface.

The interplay between ionizing radiation and life is rich with complexity and conflicting effects, and the topics highlighted here represent many fruitful areas for future research.

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Abbreviations

CR, cosmic rays; GCR, galactic cosmic rays; GRB, gamma-ray bursts; ISM, interstellar medium; pc, parsec (parallax of one arcsecond); SCR, solar cosmic rays; SEP, solar energetic particles.

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